

# The third realization of the International Celestial Reference Frame by very long baseline interferometry<sup>★</sup>

P. Charlot<sup>1</sup>, C. S. Jacobs<sup>2</sup>, D. Gordon<sup>3</sup>, S. Lambert<sup>4</sup>, A. de Witt<sup>5</sup>, J. Böhm<sup>6</sup>, A. L. Fey<sup>7</sup>, R. Heinkelmann<sup>8</sup>,  
 E. Skurikhina<sup>9</sup>, O. Titov<sup>10</sup>, E. F. Arias<sup>4</sup>, S. Bolotin<sup>3</sup>, G. Bourda<sup>1</sup>, C. Ma<sup>11,★★</sup>, Z. Malkin<sup>12,13</sup>, A. Nothnagel<sup>14,★★★</sup>,  
 D. Mayer<sup>6,★★★★</sup>, D. S. MacMillan<sup>3</sup>, T. Nilsson<sup>8,†</sup>, and R. Gaume<sup>15</sup>

<sup>1</sup> Laboratoire d'astrophysique de Bordeaux, Univ. Bordeaux, CNRS, B18N, Allée Geoffroy Saint-Hilaire, 33615 Pessac, France  
 e-mail: [patrick.charlot@u-bordeaux.fr](mailto:patrick.charlot@u-bordeaux.fr)

<sup>2</sup> Jet Propulsion Laboratory, California Institute of Technology, 4800 Oak Grove Drive, Pasadena, CA 91109-8099, USA

<sup>3</sup> NVI Inc. at NASA Goddard Space Flight Center, Code 61A.1, Greenbelt, MD 20771, USA

<sup>4</sup> SYRTE, Observatoire de Paris, Université PSL, CNRS, Sorbonne Université, LNE, 61 Av. de l'Observatoire, 75014 Paris, France

<sup>5</sup> Hartebeesthoek Radio Astronomy Observatory, PO Box 443, Krugersdorp 1740, South Africa

<sup>6</sup> Department of Geodesy and Geoinformation, Technische Universität Wien, Wiedner Hauptstraße 8-10, 1040 Vienna, Austria

<sup>7</sup> U.S. Naval Observatory, 3450 Massachusetts Avenue NW, Washington, DC 20392-5420, USA

<sup>8</sup> Helmholtz Centre Potsdam, German Research Centre for Geosciences, Telegrafenberg, A17, 14473 Potsdam, Germany

<sup>9</sup> Institute of Applied Astronomy, Russian Academy of Sciences, Nab. Kutuzova 10, St. Petersburg 191187, Russia

<sup>10</sup> Geoscience Australia, PO Box 378, Canberra, ACT 2601, Australia

<sup>11</sup> NASA Goddard Space Flight Center, Code 61A.1, Greenbelt, MD 20771, USA

<sup>12</sup> Pulkovo Observatory, St. Petersburg 196140, Russia

<sup>13</sup> Kazan Federal University, Kazan 420000, Russia

<sup>14</sup> Institut für Geodäsie und Geoinformation, Universität Bonn, Nußallee 17, 53115 Bonn, Germany

<sup>15</sup> National Science Foundation, 2415 Eisenhower Avenue, Alexandria, Virginia 22314, USA

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## ABSTRACT

A new realization of the International Celestial Reference Frame (ICRF) is presented based on the work achieved by a working group of the International Astronomical Union (IAU) mandated for this purpose. This new realization follows the initial realization of the ICRF completed in 1997 and its successor, ICRF2, adopted as a replacement in 2009. The new frame, referred to as ICRF3, is based on nearly 40 years of data acquired by very long baseline interferometry at the standard geodetic and astrometric radio frequencies (8.4 and 2.3 GHz), supplemented with data collected at higher radio frequencies (24 GHz and dual-frequency 32 and 8.4 GHz) over the past 15 years. State-of-the-art astronomical and geophysical modeling has been used to analyze these data and derive source positions. The modeling integrates, for the first time, the effect of the galactocentric acceleration of the solar system (directly estimated from the data) which, if not considered, induces significant deformation of the frame due to the data span. The new frame includes positions at 8.4 GHz for 4536 extragalactic sources. Of these, 303 sources, uniformly distributed on the sky, are identified as “defining sources” and as such serve to define the axes of the frame. Positions at 8.4 GHz are supplemented with positions at 24 GHz for 824 sources and at 32 GHz for 678 sources. In all, ICRF3 comprises 4588 sources, with three-frequency positions available for 600 of these. Source positions have been determined independently at each of the frequencies in order to preserve the underlying astrophysical content behind such positions. They are reported for epoch 2015.0 and must be propagated for observations at other epochs for the most accurate needs, accounting for the acceleration toward the Galactic center, which results in a dipolar proper motion field of amplitude 0.0058 milliarcsecond yr<sup>−1</sup> (mas yr<sup>−1</sup>). The frame is aligned onto the International Celestial Reference System to within the accuracy of ICRF2 and shows a median positional uncertainty of about 0.1 mas in right ascension and 0.2 mas in declination, with a noise floor of 0.03 mas in the individual source coordinates. A subset of 500 sources is found to have extremely accurate positions, in the range of 0.03–0.06 mas, at the traditional 8.4 GHz frequency. Comparing ICRF3 with the recently released *Gaia* Celestial Reference Frame 2 in the optical domain, there is no evidence for deformations larger than 0.03 mas between the two frames, in agreement with the ICRF3 noise level. Significant positional offsets between the three ICRF3 frequencies are detected for about 5% of the sources. Moreover, a notable fraction (22%) of the sources shows optical and radio positions that are significantly offset. There are indications that these positional offsets may be the manifestation of extended source structures. This third realization of the ICRF was adopted by the IAU at its 30th General Assembly in August 2018 and replaced the previous realization, ICRF2, on January 1, 2019.

**Key words.** reference systems – astrometry – techniques: interferometric – quasars: general – galaxies: nuclei – radio continuum: general

<sup>★</sup> Full Tables 10, 11, 12, 14, 15, and 16 are only available at the CDS via anonymous ftp to [cdsarc.u-strasbg.fr](https://cdsarc.u-strasbg.fr) (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/cat/J/A+A/644/A159>

<sup>★★</sup> Retired.

<sup>★★★</sup> Now at Technische Universität Wien, Vienna, Austria.

<sup>★★★★</sup> Now at Federal Office of Metrology and Surveying, Vienna, Austria.

<sup>†</sup> Now at Lantmäteriet – The Swedish mapping, cadastral and land registration authority, Geodetic Infrastructure, Gävle, Sweden.

## 1. Introduction

The International Celestial Reference Frame (ICRF) and its successor, ICRF2, have been the basis for high-accuracy astrometry for more than two decades. Both frames have drawn on simultaneous 8.4 GHz and 2.3 GHz observations of compact extragalactic radio sources acquired by very long baseline interferometry (VLBI), starting from the end of the 1970s. The primary frequency in this observing scheme is 8.4 GHz (*X* band), while 2.3 GHz (*S* band) is only used for ionosphere calibration. Hereafter, such dual-frequency VLBI observations are referred to as *S/X* band, following the standard designation.

The ICRF (Ma et al. 1998) was the first all-sky realization of an extragalactic frame with milliarcsecond (mas) position accuracy and the first realization of the International Celestial Reference System (ICRS) (Arias et al. 1995). It was adopted by the International Astronomical Union (IAU) at its 23rd General Assembly in 1997, replacing the Fifth Fundamental Catalog of stars (FK5) (Fricke et al. 1988) as the fundamental celestial reference frame as of January 1, 1998. Unlike previously (e.g., for the FK5 or other stellar frames), the definition of the frame axes was no longer related to the equinox and equator but relied on coordinates of so-called defining sources. The ICRF included 212 such defining sources out of a total of 608 objects for which positions were reported. The orientation of the ICRF axes was deemed to be accurate to 20 microarcseconds ( $\mu$ as) while source coordinates had a noise floor of 250  $\mu$ as. As a consequence of the ICRS definition, all source positions in the ICRF were independent of epoch. The release of ICRF was a major step forward and as such it became the required passage to link other celestial reference frames to the ICRS, among which were the dynamical frame (Standish 1998) and the Hipparcos stellar frame (Kovalevsky et al. 1997). The ICRF has also allowed for many advances in other fields, such as geodesy and Earth orientation studies, or related to practical applications like deep space navigation.

As time passed, VLBI observing networks and technology improved and further sources were observed. On the geodesy side, observations became organized in a formal way under the umbrella of the International VLBI Service for Geodesy and Astrometry (IVS), established in 1999 (Schlüter & Behrend 2007), permitting more resources to be pooled together. Two ICRF extensions, ICRF-Ext.1 and ICRF-Ext.2, were constructed in 2000 and 2002, adding another 109 sources to the frame (Fey et al. 2004). At the same time, systematic surveys of the VLBI sky were initiated using the Very Long Baseline Array (VLBA)<sup>1</sup>, enlarging the pool of VLBI sources with milliarcsecond-accurate positions to more than a thousand (Beasley et al. 2002). By the mid 2000s, it was realized that the amount of data, their accuracy, and the denser sky coverage, along with modeling improvements since ICRF was delivered, would justify the building of a new reference frame to get the full potential of the available data sets. This prompted the construction of the second realization of the ICRF, named ICRF2, which was completed in 2009 and adopted by the IAU at its 27th General Assembly in the same year (Fey et al. 2015). As a result, ICRF2 replaced ICRF on January 1, 2010. The new realization comprises 3414 sources, of which 295 are defining sources. The orientation of its axes is known to 10  $\mu$ as, while source coordinates have a noise floor of 40  $\mu$ as. The large increase in the number of sources from ICRF to ICRF2 (a factor of six) was largely due to the inclusion of observations

from a series of VLBA Calibrator Survey (VCS) astrometric campaigns carried out between 1994 and 2007 (Beasley et al. 2002; Fomalont et al. 2003; Petrov et al. 2005, 2006, 2008; Kovalev et al. 2007). The goal of these was to expand the pool of calibrators available for VLBI observations in phase-referencing mode (Beasley & Conway 1995). Such campaigns added nearly 2200 sources exclusively observed by the VLBA to the catalog of sources derived from IVS sessions. The VCS sources, still, had position uncertainties typically five times larger than the other sources due to being observed in survey mode and generally only in a single session. For this reason, they were categorized separately from the IVS sources in ICRF2.

Since the release of ICRF2, the VLBI database has continued to expand thanks to ongoing observing programs run by the IVS but also through specific projects carried out independently, notably by using the VLBA. The latter includes a complete re-observation of all VCS sources in 2014–2015, which has brought an overall factor of five improvement in coordinate uncertainties, hence bringing position uncertainties for the VCS sources closer to those for the rest of the ICRF2 sources (Gordon et al. 2016). A specific effort has also been made to strengthen observations of optically bright ICRF2 sources within IVS programs to facilitate the alignment of the optical reference frame which is being built by the *Gaia* space mission (Le Bail et al. 2016). At the same time, VLBI observations at higher radio frequencies began to develop, namely at 24 GHz (*K* band) and 32 GHz (*Ka* band), the latter with simultaneous *X* band observations for ionosphere calibration (hence the usual *X/Ka* band designation for this dual-frequency observing scheme). An initial catalog of 268 sources, together with VLBI images of the sources, was produced using the VLBA at *K* band (Lanyi et al. 2010; Charlot et al. 2010), while positions for 482 sources were reported at *X/Ka* band from observations with the Deep Space Network (DSN) (Jacobs et al. 2012). Despite the limited data sets, both catalogs showed an overall agreement with ICRF2 at the 300  $\mu$ as level when comparing individual source coordinates, hence revealing the value of such high-frequency observations. By 2012, the wealth of the additional VLBI data already acquired, or foreseen, together with the need to have a state-of-the-art VLBI frame to align as well as possible the future *Gaia* optical frame onto the ICRS, created the necessary motivation for generating a new realization of the ICRF. Shortly after the 28th General Assembly of the IAU in Beijing, a working group under IAU Division A was assembled to this end. The mandate of the working group was to generate the third realization of the ICRF by 2018, for adoption at the 30th IAU General Assembly to take place that year.

The work accomplished toward the generation of the third realization of the ICRF, hereafter referred to as ICRF3, was organized along several lines. One such line was aimed at acquiring new appropriate data to correct deficiencies of ICRF2. In this respect, the focus was placed not so much on trying to increase the number of sources but rather on improving uniformity and internal consistency. The campaign to re-observe all VCS sources (Gordon et al. 2016) falls into this line. As noted above, this campaign vastly improved coordinate uncertainties for the relevant 2200 such sources, resulting in an overall distribution of source position uncertainties that is now much more uniform. Specific efforts were also targeted to strengthen observations in the far south (i.e., for declinations below  $-45^\circ$ ). However, the limited number of VLBI telescopes in the Southern Hemisphere remains as a major bottleneck to reach the same quality in that area of the sky (in terms of source density and position accuracy) as that available further north. The full data set used to generate ICRF3 is described in Sect. 2. Compared to

<sup>1</sup> The VLBA is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc.

ICRF and ICRF2, a new feature is the inclusion of data at  $K$  band and  $X/Ka$  band in addition to those at the standard  $S/X$  band geodetic frequencies. The resulting three-frequency positions (at  $X$ ,  $K$ , and  $Ka$  band) are herewith reported as part of ICRF3 without being combined in order to preserve the underlying astrophysical information.

Another major activity of the working group consisted in generating ICRF3 prototype realizations at several stages of the work. These allowed the group to study the impact of data sets, astronomical and geophysical modeling, analysis configuration, software packages, and individuals that analyze the data on the resulting frame. In practice, such prototype realizations were produced at eight different institutions<sup>2</sup> in Australia, Austria, France, Germany, Russia, and USA using five different software packages. Of interest is that one such ICRF3 prototype realization was delivered to the *Gaia* science team in July 2017 and used in the process to generate the catalog for the *Gaia* Data Release 2 (*Gaia* DR2) in order to align the *Gaia* frame with the ICRS (Lindgren et al. 2018). The final solution for ICRF3 incorporates data up to spring 2018 and was produced in July 2018. Section 3 below reviews the adopted modeling and analysis configuration, while Sect. 4 describes the alternate analyses that have been conducted to assess errors in ICRF3. A newly added feature in the modeling is the galactocentric acceleration of the solar system, long sought and first detected by Titov et al. (2011). With a magnitude of about  $5 \mu\text{as yr}^{-1}$ , this effect produces detectable apparent drifts in the source positions, especially over the almost 40-year span now reached by the  $S/X$  band data. Unlike previously (i.e., for ICRF and ICRF2), the source coordinates in ICRF3 are referred to a specific epoch and hence should be properly propagated for epochs away from that reference epoch, accounting for Galactic acceleration, for the most demanding needs.

Details of ICRF3 are given in Sect. 5, including source categorization and tables of positions reported separately at  $X$  band,  $K$  band, and  $Ka$  band. The frame has a new set of defining sources selected in a way to be uniformly distributed on the sky. Recommendations to future users on how to use the cataloged ICRF3 positions, depending on their needs, are also provided. Section 6 makes an assessment of the alignment of ICRF3 onto ICRS, emphasizes the improvement and benefits of the new realization over ICRF2, and compares ICRF3 with the *Gaia* DR2 celestial reference frame (*Gaia*-CRF2), which is the first extragalactic frame ever built in the optical domain (Gaia Collaboration 2018). Consistency between multi-frequency radio positions and between radio and optical positions is also addressed as part of that section. The final sections outline the adoption process by the IAU and the future evolution of the ICRF.

## 2. Observations

The VLBI data used to build ICRF3 were acquired by arrays of 2–20 radiotelescopes organized in their vast majority under the umbrella of the IVS, the VLBA, and the DSN. Over the years, a total of 167 telescopes, located on 126 different sites, participated in such VLBI sessions (Fig. 1). Observations were carried out using the so-called bandwidth synthesis mode, which

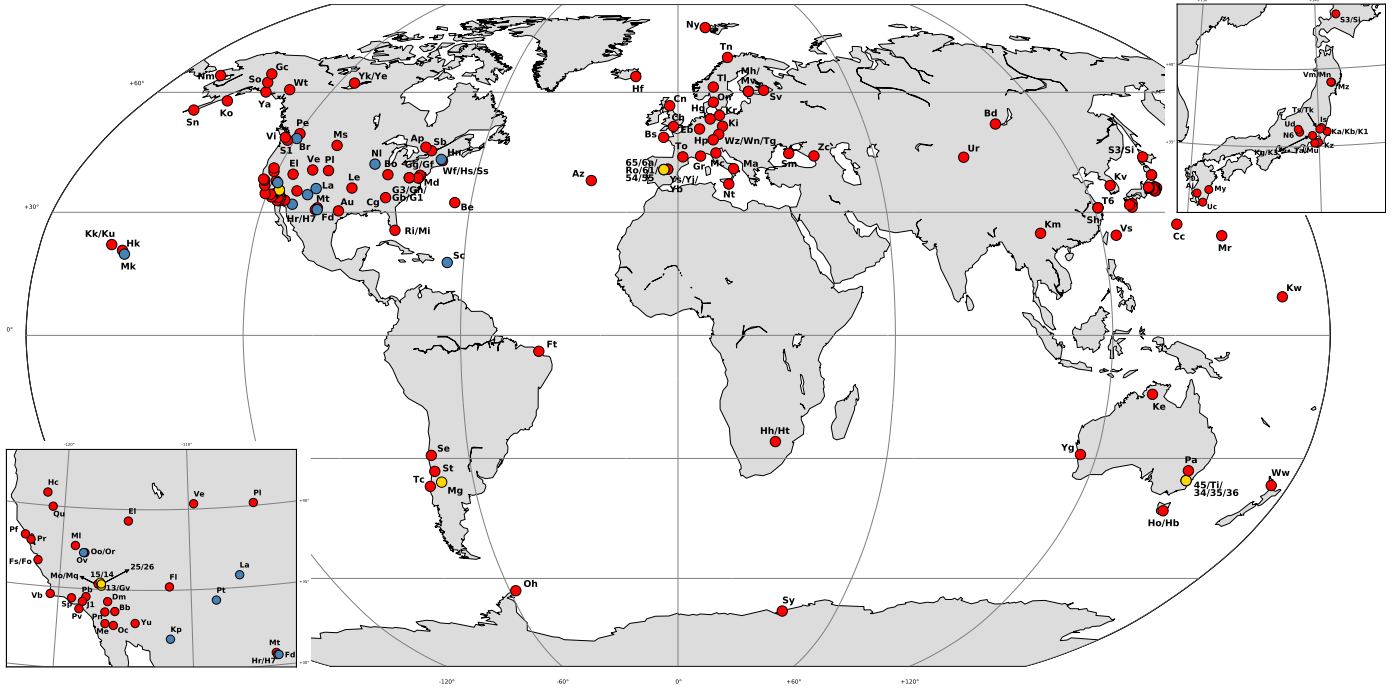
permits the determination of precise group delay quantities by observing multiple channels spread out across a bandwidth of several hundred MHz, as originally devised by Rogers (1970). Following acquisition, the data were processed at one of the IVS correlators (in Bonn, Haystack, Kashima, Shanghai, Vienna, or Washington), the VLBA correlator in Socorro, the Australia Telescope National Facility correlator in Perth, or the DSN processor in Pasadena. Post-processing was accomplished by calibrating the raw phases to make them consistent in all channels and by fringe-fitting these to obtain the group delay quantities which are further fed into geodetic and astrometric software packages for the estimation of source positions (see Sect. 3 below). Such post-processing was conducted either at the correlators or at the institutions coordinating the relevant VLBI sessions or sets of sessions. Dual-frequency observing at  $S/X$  band has been standard since the early days of geodetic and astrometric VLBI as it permits the calibration of the dispersive delay caused by the ionosphere using a combination of the measurements at the two frequencies. As noted above, a similar scheme was implemented for observing at  $Ka$  band (with simultaneous measurements at  $X$  band), while at  $K$  band observing has remained single-frequency, hence requiring proper modeling of the ionospheric delays for the latter. In general, VLBI sessions are 24-hour long in order to separate parameters for polar motion and nutation and to average out unmodeled geophysical effects which vary on a diurnal basis. Each session generally observes a few tens to a few hundreds of sources depending on the size of the network, the slewing speed of the antennas, the data recording rate, and the objective of the session (i.e., whether it is a survey program). The number of observations collected during a session varies, depending in the first place on the size of the network. Over the years, the amount of data acquired has increased, due to larger networks being used, culminating in 2017 with more than one million observations collected at  $S/X$  band, 0.1 million collected at  $K$  band, and 0.01 million collected at  $X/Ka$  band (see the distribution of observations per year in Fig. 2). Characteristics of the data sets at each of the three frequency bands are given in the subsections below.

### 2.1. $S/X$ band (2.3/8.4 GHz)

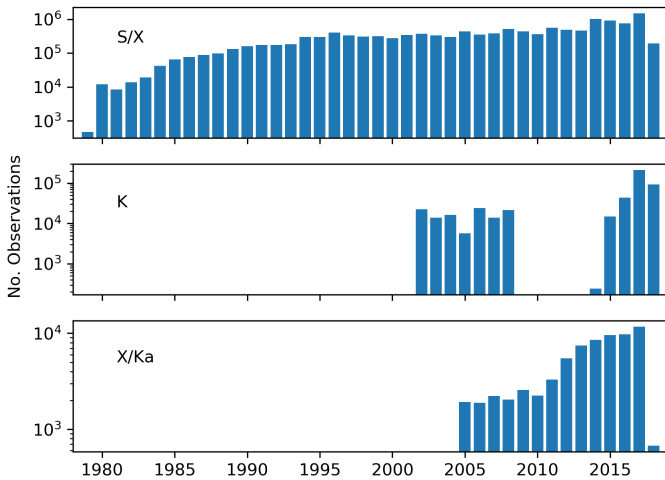
The data used for ICRF3 in this frequency band include the entire pool of geodetic and astrometric VLBI sessions acquired and made available by the IVS with rare exceptions. The sessions range from August 3, 1979 to March 27, 2018 and come from a variety of programs dedicated to monitor the Earth orientation, establish the terrestrial and celestial reference frames, and maintain and expand these, according to the objectives of the IVS (Schuh & Behrend 2012). Details on the current observing programs are given in Nothnagel et al. (2017). Data acquired prior to the establishment of the IVS in 1999 had similar goals and were obtained through ad hoc arrays organized by cooperations between individual observatories or national agencies and are described in Ma et al. (1998). Of particular interest for the celestial frame are the Research and Development VLBA (RDV) sessions conducted jointly with the VLBA six times a year (since 1997) which assemble a network of 15–20 stations, allowing for observation of 80–100 sources each time. Such sessions are also essential to image the sources and assess their suitability as fiducial points to materialize the celestial frame (see Sect. 5.2). Southern-Hemisphere IVS sessions are equally important, even if conducted with arrays of only a few stations, as they allow for observations in the far south which otherwise would be impossible with the northern arrays. Unfortunately,

<sup>2</sup> Geoscience Australia (Australia), Technische Universität Wien (Austria), Observatoire de Paris (France), Helmholtz Centre Potsdam (Germany), Institute of Applied Astronomy St. Petersburg (Russia), NASA Goddard Space Flight Center (USA), Jet Propulsion Laboratory (USA), U.S. Naval Observatory (USA).





**Fig. 1.** World map showing the geographical location of the 167 antennas (situated on 126 different sites) that participated in the observations used for ICRF3. The red dots show the antennas from the IVS network (and pre-existing adhoc VLBI arrays that observed at  $S/X$  band), the blue ones those from the VLBA, and the yellow ones those from the DSN and ESA. The two-character codes printed near each dot correspond to the short names of the antennas, as defined in the IVS nomenclature. The two insets show enlargements of western US and Japan where a large number of antennas (including mobile VLBI stations) have been used to collect geodetic VLBI data over the years due to the seismic nature of these regions.



**Fig. 2.** Distribution of the observations used for ICRF3. The three histograms show the number of VLBI delays per year at  $S/X$  band (upper panel),  $K$  band (middle panel), and  $X/Ka$  band (lower panel). For ease of reading, the number of observations is plotted with a logarithmic scale.

despite ongoing efforts (Plank et al. 2017), the paucity of observations in this area of the sky compared to the north remains a deficiency in the data set. Occasional astrometric observations conducted by other VLBI arrays such as the European VLBI Network and the Australian Long Baseline Array have also been incorporated.

The generation of ICRF3, furthermore, took advantage of a number of dedicated astrometric sessions conducted with the VLBA at  $S/X$  band since the mid 1990s. These include the series of VCS campaigns that took place between 1994 and 2007,

already incorporated in ICRF2 (see Fey et al. 2015), along with the campaign that re-observed all VCS sources in 2014–2015, which was initiated specifically for the purpose of ICRF3, as reported in Gordon et al. (2016). More recently, another 24 such VLBA sessions have been run under the US Naval Observatory share of the VLBA observing time. Those sessions targeted all sources in the VLBI pool meeting one of the following criteria: (i) had not been observed since 2009 (i.e., when ICRF2 was delivered), (ii) had less than 50 observations, (iii) were observed in three or fewer sessions, or (iv) were among the weakest known optically bright sources listed in Le Bail et al. (2016). The goal here was to further enhance uniformity of the data sets for ICRF3. In all, the VLBA sessions constitute only a small portion of the entire set of  $S/X$  band sessions (200 sessions out of a total of 6206 sessions) but account for 26% of the data, while more than two-thirds (68%) of the sources have observations coming exclusively from the VLBA. See Gordon (2017) for further details on the impact of the VLBA observing on the celestial frame.

## 2.2. $K$ band (24 GHz)

The data sets in this band are made of 40 VLBA sessions that observed the northern sky (down to mid-southern declinations), supplemented with 16 single-baseline sessions between telescopes in Hartebeesthoek (South Africa) and Hobart (Australia) that observed sources below  $-15^\circ$  declination (down to the far south). The first VLBA session was conducted on May 15, 2002 and was part of a set of ten such sessions that ultimately led to the first realization of a celestial frame in this frequency band, although not covering the entire sky (Lanyi et al. 2010). Apart from two similar follow-up sessions in 2008, VLBA observing was then interrupted until 2015, after which it was started again

(see Fig. 2). Since then, another 25 VLBA sessions have been carried out, the bulk of which were run in 2017–2018 under the US Naval Observatory share of the VLBA observing time. The latest session incorporated in this work was conducted on May 5, 2018. Additionally, three archived VLBA sessions dedicated to observing sources in the Galactic plane (Petrov et al. 2011) were also included. Observations between Hartebeesthoek and Hobart were initiated at about the same time as the VLBA sessions restarted (first session run on May 4, 2014) to complete the sky coverage in the far south. One southern session also included the Tianma 65 m telescope near Shanghai (China) while another one included the Tidbinbilla 70 m telescope near Canberra (Australia). In all, the VLBA sessions account for 99% of the data while the Southern-Hemisphere sessions account for 1% of it, which leads to a majority of sources (66%) having only VLBA observations, a similar division as that at *S/X* band.

### 2.3. *X/Ka* band (8.4/32 GHz)

Observations at *X/Ka* band were initiated in 2005 with the primary goal of building a reference frame for spacecraft navigation, now conducted on the DSN using the *Ka* frequency band. The data set includes a total of 168 single-baseline sessions that involved seven telescopes at the three DSN sites in Goldstone (California), Robledo (Spain), and Tidbinbilla (Australia). The first of these sessions took place on July 9, 2005 while the most recent one incorporated was run on January 28, 2018. Occasionally (in about 10% of the sessions), the European Space Agency (ESA) telescope in Malargüe (Argentina) joined the observations, which was essential to improve the north-south geometry of the network and reduce systematics in the reference frame.

## 3. Analysis

The principle behind VLBI data analysis is to compare measured quantities with a priori theoretical modeling of the same quantities and to refine underlying models by estimating model parameter corrections that best fit the data. Depending on the objective pursued, such parameters may pertain to the entire data set (e.g., station positions and velocities, source positions,...) or only to individual sessions (e.g., the Earth orientation parameters,...) or even to a small portion of a session (e.g., clock and troposphere parameters which vary on the order of hours). Least-squares methods are generally employed for this estimation. The VLBI modeling, analysis configuration, and software packages used to produce ICRF3 are outlined in the following sections.

### 3.1. Astronomical, geophysical, and instrumental modeling

The measured VLBI quantities (group delay and delay rate) used for ICRF3 were analyzed employing state-of-the-art astronomical and geophysical modeling, generally following the prescriptions of the International Earth Rotation and Reference Systems Service (IERS) (Petit & Luzum 2010). An extensive review of all the effects to be incorporated in the VLBI model in order to reach the highest accuracy is given in Sovers et al. (1998). Apart from the effect induced by the galactocentric acceleration of the solar system, only brief information, primarily to identify the models selected, is thus provided here. The interested reader is referred to the Sovers et al. (1998) review for details on the underlying physics. Galactocentric acceleration is addressed specifically in Sect. 3.2 as it is the first time this effect is introduced in the modeling used to generate a VLBI celestial reference frame.

The geometric portion of the VLBI delay, including relativistic effects, was derived consistently with the so-called consensus model (Eubanks 1991) which provides 1 ps accuracy. Formulation of the geometric VLBI delay necessitates describing completely the evolution of the dynamic Earth over the period of the observations. This requires specifying the orientation of the Earth's spin axis in inertial space (i.e., precession and nutation), and relative to the Earth's crust (polar motion), and to characterize the daily rotation of the Earth around that axis (UT1). Also to be considered are the various deformations that affect the Earth's crust on which the radiotelescopes are attached. These comprise tectonic plate motions, tidal deformations, and atmospheric pressure loading effects. Modeling of the Earth's spin axis was achieved using the MHB nutation (Mathews et al. 2002) and P03 precession (Capitaine et al. 2003; Hilton et al. 2006), further designated as IAU 2000A nutation and IAU 2006 precession after adoption of these models by the IAU. A priori polar motion and UT1 were retrieved from the IERS Rapid Service/Prediction Centre (solution labeled “finals.data”, see Dick & Thaller 2018, Sect. 3.5.2), to which were added short-period tidal variations, as prescribed by the IERS (see Petit & Luzum 2010, Chap. 8). Initial station (radiotelescope) positions and velocities were taken from the ITRF2014 terrestrial frame (Altamimi et al. 2016), incorporating post-seismic deformation models for sites that were subject to major earthquakes, and further adding deformations due to solid Earth tides, ocean loading, and atmospheric pressure loading. Displacements caused by solid Earth tides were derived following the IERS prescriptions (see Petit & Luzum 2010, Sect. 7.1.1). Ocean loading displacements were obtained from the TPXO.7.2 model (Egbert & Erofeeva 2002), supplemented with the FES99 model (Lefèvre et al. 2002) for the long-period Ssa tide, while those due to atmospheric pressure loading (both tidal and non-tidal) come from the APLO model (Petrov & Boy 2004). Further displacements caused by the centrifugal effect of polar motion on the solid Earth (see Petit & Luzum 2010, Sect. 7.1.4) and the oceans (Desai 2002) were also incorporated in the modeling. Calculation of the geometric VLBI delay finally considered a component resulting from the thermal expansion of the antennas which are subject to structural deformations when temperature varies, hence causing a displacement of the position of the reference point of the instruments (Nothnagel 2009).

Added to the geometric delay were corrections for atmospheric propagation, including the contribution due to high-altitude charged particles (ionosphere) and that due to the neutral component (troposphere). Ionospheric delays are proportional to the inverse of the frequency-squared and were calibrated using the dual-frequency data collected at *S/X* band and *X/Ka* band, whereas they were modeled using total electron content maps for *K* band. Such ionospheric maps are produced daily from global navigation satellite systems and were retrieved from NASA's Crustal Dynamics Data Information System (Noll 2010) for the purpose of the *K* band analysis. Tropospheric delays were derived using the VMF1 mapping function (Böhm et al. 2006) to map zenith delay contributions to relevant elevations. It must be noted that observations below 5° elevation were discarded due to inadequacies in modeling the troposphere at such low elevations. The zenith delays themselves were estimated every 30 min at each site using a continuous piecewise linear function during the least-squares parameter adjustment. Additional east-west and north-south tropospheric gradients were also estimated. The interval between such estimates was six hours for the *S/X* band and *K* band analyses, whereas they were estimated only once for the entire data set at *X/Ka* band due to the more

limited sky coverage above each site for the latter. Corrections for instrumental delays were further added to the above geometric and atmospheric VLBI delay components. These come from propagation delays in the cables (which have a different length at each telescope), lack of synchronization of the clocks between stations, and clock instabilities. In practice, since it is not possible to calibrate those instrumental delays precisely, they were treated altogether, assuming an overall clock-like behavior at each telescope. A 60 min continuous piecewise linear function, with quadratic terms when needed, was used to model them accordingly, the parameters of which were estimated during the least-squares parameter adjustment, as in the case of troposphere.

### 3.2. Galactocentric acceleration

The relative motion of the solar system barycenter with respect to the extragalactic reference frame may cause apparent changes in the radio source positions due to aberrational effects. In this motion, only the non-linear part (i.e., the acceleration) is to be considered since the constant part is absorbed into the reported source positions by convention. As noted in [Sovers et al. \(1998\)](#), the said motion is conveniently decomposed into three components: (i) the motion of the solar system barycenter with respect to the Galactic center, (ii) the motion of the Galaxy relative to the Local Group, and (iii) the motion of the Local Group relative to the extragalactic frame (which may be assumed at rest relative to the cosmic microwave background since the observed radio sources are located at cosmological distances). Of these components, only the first is expected to induce significant acceleration on a timescale of a few decades such as that of our VLBI data set. It is thus solely considered in the rest of this section.

Based on the rotational property of the Galaxy, the solar system barycenter is known to move around the Galactic center with an orbital period of about 200 million years. Assuming a purely circular rotation, the resulting acceleration vector, which reflects the curvature of the orbit, is directed toward the Galactic center. Classically, this acceleration translates into an aberration effect which takes the form of an apparent overall proper motion for the distant extragalactic radio sources. The amplitude  $A_G$  of this proper motion (see, e.g., [Kovalevsky 2003](#)) is given by

$$A_G = \frac{V_0^2}{R_0 c}, \quad (1)$$

where  $R_0$  is the distance of the barycenter of the solar system to the Galactic center,  $V_0$  is the linear speed along its orbit, and  $c$  is the speed of light. As also pointed out by [Kovalevsky \(2003\)](#), the resulting effect varies according to the region of the sky since it depends on the projection of the acceleration vector (which points from the solar system barycenter to the Galactic center) onto the plane of the sky in the source direction. In practice, it is conveniently mapped using Galactic coordinates. For a source at Galactic longitude  $l$  and Galactic latitude  $b$ , the components of the corresponding proper motion,  $\mu_l \cos b$  and  $\mu_b$ , are then

$$\mu_l \cos b = -A_G \sin l, \quad (2)$$

$$\mu_b = -A_G \sin b \cos l. \quad (3)$$

Going further and adopting the IAU-recommended values for the Galactic constants<sup>3</sup>,  $R_0 = 8.5$  kpc and  $V_0 = 220$  km s<sup>-1</sup>, a proper motion amplitude  $A_G$  of  $4 \mu\text{as yr}^{-1}$  is obtained from Eq. (1). With

<sup>3</sup> Recent estimates of  $R_0$  and  $V_0$  deviate somewhat from these IAU-recommended values which date back to 1985. See [Vallée \(2017\)](#).

four decades of accumulated VLBI data, apparent source displacements of up to  $150 \mu\text{as}$  are thus expected, which is quite significant considering current source position accuracies of a few tens of microarcseconds (see below). Galactic acceleration therefore clearly needs to be considered in the modeling.

The galactocentric acceleration of the solar system was predicted to induce detectable proper motions for the extragalactic sources soon after the inception of the VLBI astrometric technique ([Fanselow 1983](#)). Attempts to detect such proper motions from the VLBI data date back to the 1990s, taking advantage of data sets that were already more than a decade long ([Sovers & Jacobs 1996](#); [Gwinn et al. 1997](#)). At the same time, plans started to develop to measure those proper motions from future optical space astrometry ([Bastian 1995](#); [Mignard 2002](#); [Kovalevsky 2003](#); [Kopeikin & Makarov 2006](#)). The actual detection of the effect was made by [Titov et al. \(2011\)](#) who estimated an amplitude of  $6.4 \pm 1.5 \mu\text{as yr}^{-1}$ , in reasonable agreement with the above prediction. This result was derived through a vector spherical harmonics analysis of time series of VLBI source coordinates covering two decades. Several other determinations followed, based on increasing VLBI data span and/or different analysis schemes, including the estimation of the Galactic acceleration amplitude as part of a global VLBI solution ([Xu et al. 2012](#); [Titov & Lambert 2013](#); [Titov & Krásná 2018](#)). In all, the values derived range from  $5.2 \mu\text{as yr}^{-1}$  to  $6.4 \mu\text{as yr}^{-1}$  with an uncertainty of  $0.3 \mu\text{as yr}^{-1}$  for the most recent determinations. Additional estimates can be inferred from measurements of parallax and proper motions of Galactic masers using VLBI phase-referenced techniques ([Reid et al. 2009, 2014](#); [Brunthaler et al. 2011](#); [Honma et al. 2012](#)). These point to somewhat lower values, from  $4.8$  to  $5.4 \mu\text{as yr}^{-1}$ , with similar uncertainties as the geodetic VLBI determinations. See [MacMillan et al. \(2019\)](#) for an overview of all such determinations either from geodetic VLBI or from maser proper motions. An open question is whether the acceleration vector of the solar system barycenter is offset from the Galactic center, which would happen if the Sun were subject to a specific peculiar motion apart from its circular rotation around the Galactic center. When estimated from geodetic VLBI data, values of this offset range from non-significant (i.e., near  $0^\circ$ ) to about  $20^\circ$ , with uncertainties of  $5$ – $10^\circ$  ([Titov et al. 2011](#); [Xu et al. 2012](#); [Titov & Lambert 2013](#); [Titov & Krásná 2018](#)). In all, one cannot be sure whether there is an actual offset or whether some of the above estimates just reflect systematic errors in the data.

For the present work, we decided to re-determine the amplitude of the Galactic acceleration because we saw no compelling evidence to adopt one or the other of the published values (nor any of the unpublished values known to us) at the time. Additionally, the data sets on which those previous determinations are based are only up to 2016 whereas the data sets for ICRF3 extend to 2018. Because of the high-quality data acquired in the period, we thought these additional two years could make a difference. We also thought that it was most appropriate to use a value determined from a data set that covers the same time span as that for ICRF3. A dedicated analysis estimating both the amplitude and direction of the acceleration vector of the solar system barycenter was thus performed, prior to constructing ICRF3, based on the almost 40-year long  $S/X$  data set in our hand. While this estimation, ideally, could have been accomplished simultaneously with the determination of the  $S/X$  frame, we decided not to do so because the analysis configuration adopted for ICRF3 treats all sources as global parameters (see Sect. 3.3 below) and this scheme is not entirely appropriate for estimating the solar system barycenter acceleration vector. As



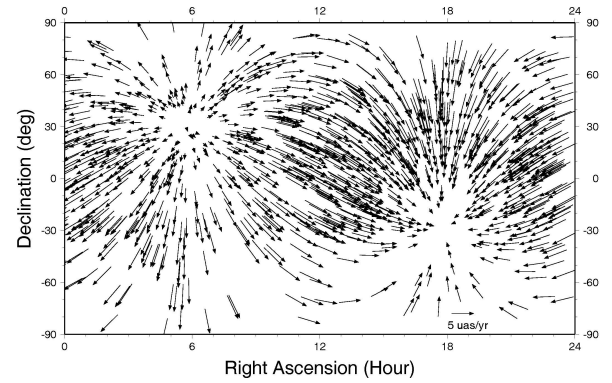
remarked by Titov et al. (2011), some sources, mostly observed in the early VLBI sessions, are subject to notable instabilities (due to their having extended and variable structures) which affect significantly the estimation of the acceleration parameters if not filtered out in the analysis process. In the realization of ICRF2, those sources were treated as arc parameters (i.e., a new position was estimated for each session) and denoted as special handling sources (Fey et al. 2015). For our determination, we adopted a similar approach, meaning that the positions for the 39 sources identified as such in ICRF2 were estimated separately for each session in which they were observed in order to limit the impact of their positional variability on the derived acceleration parameters.

The dedicated analysis described in the previous paragraph led to a value of  $5.83 \pm 0.23 \mu\text{as yr}^{-1}$  for the amplitude of the solar system barycenter acceleration vector, while the estimated vector was found to point in the direction  $\alpha_{\text{SSA}} = 270.2^\circ \pm 2.3^\circ$ ,  $\delta_{\text{SSA}} = -20.2^\circ \pm 3.6^\circ$ . This direction is within  $10^\circ$  of the Galactic center (located at  $\alpha_G = 266.4^\circ$ ,  $\delta_G = -29.0^\circ$ ) and the significance of the offset is less than  $2.5\sigma$ . Conversely, the amplitude of the acceleration vector is detected at the  $25\sigma$  level. Adopting a conservative approach, we decided not to consider any such offset in the construction of ICRF3, in view of its marginal significance, and therefore assumed that the solar system barycenter acceleration vector points to the Galactic center. Concerning the amplitude of the acceleration vector, we adopted the value inferred from our analysis, namely  $5.8 \mu\text{as yr}^{-1}$ . This value was then applied in the modeling for the analysis of all observations considered for ICRF3. The apparent proper motion field induced over the celestial sphere by this correction is depicted in Fig. 3. Another decision was concerned with the reference epoch of the source coordinates to report in ICRF3 since these are no longer invariant with time due to the application of the acceleration correction. While a perhaps natural choice would have been 2000.0, this date is now more than 20 years back and would have made the acceleration corrections quasi mandatory for use of the frame in present times. We thus decided to leave out that option and adopted instead 2015.0 as reference epoch. Apart from being closer to the present day, this date is also close to the mean epoch of the observations for ICRF3 (2014.5 at *S/X* band, 2016.7 at *K* band, and 2015.3 at *X/Ka* band). Unlike epoch 2000.0, such a choice facilitates the use of the frame in present times where corrections for Galactic acceleration will remain small and may not be necessary unless the highest position accuracy is required. Additionally, it is also very close to the reference epoch of *Gaia* DR2, which is J2015.5 (Lindgren et al. 2018; Gaia Collaboration 2018), hence making comparisons between the two frames more straightforward. Finally, it is to be underlined that all our conclusions have been fed into the IVS working group on Galactic aberration which reported its work in MacMillan et al. (2019).

### 3.3. Configuration of analysis

Aside from the astronomical, geophysical, and instrumental modeling described above, the remaining elements of the data analysis to configure were concerned with the selection of parameters to estimate, the mechanisms to define the orientation of the terrestrial and celestial frames, and the data weighting scheme. Each of these aspects is discussed in turn below.

Besides clocks and tropospheric parameters, which must be solved for since they are a source of nuisance and cannot be modeled precisely as discussed in Sect. 3.1, the other parameters that were estimated in the analysis include station positions and



**Fig. 3.** Aberration proper motion field (in equatorial coordinates) resulting from a solar system barycenter acceleration vector with amplitude  $5.8 \mu\text{as yr}^{-1}$ , pointing to the Galactic center. The scale is given by the length of the vector bar plotted in the lower right-hand of the figure.

velocities, the Earth Orientation Parameters (EOP), and radio source coordinates. Station positions and velocities were estimated globally from the entire data set at each frequency band, resulting in a single position and velocity estimate for each antenna (except for known discontinuities, e.g., due to earthquakes or mechanical movement of the antenna). Radio source positions were treated using a similar scheme, with a single position estimated for each source at each frequency band. No sources were made special cases, unlike in the ICRF2 analysis where 39 sources with significant position instabilities were estimated session-wise (Fey et al. 2015). We did not repeat this approach because there was no indication in the present case, based on the tests we carried out (see Sect. 4.1 below), that such sources would degrade the frame if solved globally (which was a concern for ICRF2). We thus saw no reason to treat them differently than the other sources. Adopting the same approach for all sources also guarantees consistency in the resulting source position uncertainties. Unlike antenna and source parameters, the EOP were estimated session-wise, with the exception of single-baseline sessions where they were held fixed to their a priori values. Those parameters include offsets and rates for UT1 and the two components of polar motion, and two nutation offsets, all of which were estimated at the midpoint of each session. The EOP define the rotation between the terrestrial and celestial frames and allow for the connection the two frames at any epoch.

The data analysis was configured in such a way that the terrestrial frame derived from estimating station positions and velocities is consistent with the most recent realization of the International Terrestrial Reference Frame (ITRF), namely ITRF2014 (Altamimi et al. 2016). To this end, loose no-net-translation and no-net-rotation constraints were applied to the positions and velocities of a set of core antennas that do not show any position discontinuities. This scheme ensures that the resulting frame is aligned onto the a priori frame (i.e., ITRF2014) and that their origins coincide. This set of core antennas was comprised of the 38 stations in Table 1 for the *S/X* band data set, while at *K* band it was comprised of all VLBA stations except the MK-VLBA antenna. That station had to be excluded from the constraint because it suffered an earthquake on June 15, 2006 and shows a position discontinuity at that epoch. The *X/Ka* band data set was treated differently because it consists mostly of single-baseline sessions where the EOP are kept fixed. In that case, the derived frame is implicitly consistent with the a priori polar motion and UT1 series (which are themselves consistent with ITRF2014).

**Table 1.** Name, two-character code (as in Fig. 1), and geographical location of the 38 core antennas used to align the terrestrial reference frame onto ITRF2014 in the *S/X* band analysis.

Code	Station name	Longitude	Latitude	Location
Kk	KOKEE	−159.67	22.13	USA
Ku	KAUAI	−159.67	22.13	USA
Hc	HATCREEK	−121.47	40.82	USA
Vb	VNDNBERG	−120.62	34.56	USA
Br <sup>(†)</sup>	BR-VLBA	−119.68	48.13	USA
Oo	OVRO_130	−118.28	37.23	USA
Ov <sup>(†)</sup>	OV-VLBA	−118.28	37.23	USA
Kp <sup>(†)</sup>	KP-VLBA	−111.61	31.96	USA
Pt <sup>(†)</sup>	PIETOWN	−108.12	34.30	USA
La <sup>(†)</sup>	LA-VLBA	−106.25	35.78	USA
Fd <sup>(†)</sup>	FD-VLBA	−103.94	30.64	USA
Nl <sup>(†)</sup>	NL-VLBA	−91.57	41.77	USA
Ri	RICHMOND	−80.38	25.61	USA
G3	NRAO85_3	−79.84	38.43	USA
Gn	NRAO20	−79.83	38.44	USA
Ap	ALGOPARK	−78.07	45.96	Canada
Hn <sup>(†)</sup>	HN-VLBA	−71.99	42.93	USA
Hs	HAYSTACK	−71.49	42.62	USA
Wf	WESTFORD	−71.49	42.61	USA
St	SANTIA12	−70.67	−33.15	Chile
Sc <sup>(†)</sup>	SC-VLBA	−64.58	17.76	USA
Ft	FORTLEZA	−38.43	−3.88	Brazil
Ys	YEBES40M	−3.09	40.52	Spain
Ny	NYALES20	11.87	78.93	Norway
On	ONSALA60	11.93	57.40	Sweden
Wz	WETTZELL	12.88	49.15	Germany
Nt	NOTO	14.99	36.88	Italy
Ma	MATERA	16.70	40.65	Italy
Hh	HARTRAO	27.69	−25.89	South Africa
Sv	SVETLOE	29.78	60.53	Russia
Yg	YARRA12M	115.35	−29.05	Australia
Sh	SESHAN25	121.20	31.10	China
Ke	KATH12M	132.15	−14.38	Australia
Ka	KASHIMA	140.66	35.95	Japan
Hb	HOBART12	147.44	−42.80	Australia
Ho	HOBART26	147.44	−42.80	Australia
45	DSS45	148.98	−35.40	Australia
Ww	WARK12M	174.83	−36.60	New Zealand

**Notes.** <sup>(†)</sup>Subset of antennas used to align the terrestrial frame at *K* band.

A similar approach was used to ensure that the celestial reference frames produced at *S/X* band, *K* band, and *X/Ka* band are all aligned and consistent with the previous realization of the ICRF, namely ICRF2 (Fey et al. 2015). This overall alignment was achieved in the first stage by aligning the *S/X* band frame onto ICRF2 and in the second stage by aligning the *K* band and *X/Ka* band frames onto the *S/X* band frame. This two-stage approach allowed us to take advantage of the improved *S/X* band frame (compared to ICRF2) for the alignment of the *K* band and *X/Ka* band frames. In practice, the alignment of the *S/X* band frame was accomplished by applying a tight ( $10\mu\text{as yr}^{-1}$ ) no-net-rotation constraint to the positions of the 295 ICRF2 defining sources. As discussed in Sect. 4.1 below, alternate analysis solutions, where a fraction of the ICRF2 defining sources (those that show extended structures) were left out from the constraint, were tried, but the impact on the definition of axes was found to be

minimal. Therefore, we stuck to the original 295 ICRF2 defining sources for fixing the orientation of the *S/X* band frame through that constraint. As noted above, the *S/X* band frame, once realized, then served as a reference on which to align the *K* band and *X/Ka* band frames. To this end, a no-net-rotation constraint was applied to the subset of ICRF3 defining sources included in the *K* band frame and the same was accomplished for the *X/Ka* band frame (see Sect. 5.2 below for details on the selection of the ICRF3 defining sources). In all, there were 187 usable ICRF3 defining sources at *K* band and 174 such sources at *X/Ka* band. At *K* band, six ICRF3 defining sources included in the frame were deemed to be unsuitable for use in the rotation constraints in this band because of too few ( $<10$ ) observations. In the case of *X/Ka* band, two ICRF3 defining sources included in the frame (0346+800 and 0743−006) were also left out from the rotation constraints as possible outliers. The sources used to orient the *K* band and *X/Ka* frames are those labeled as defining sources in Tables 11 and 12 below with the above restrictions.

A final aspect of the analysis configuration is the weighting of the individual measurements. Following the usual VLBI practice, the weighting factor  $w_i$  assigned to a given observation  $i$  was determined as a function of the unit weight  $\sigma_0$  as

$$w_i = \frac{\sigma_0^2}{\sigma_i^2 + \sigma_s^2}, \quad (4)$$

where  $\sigma_i$  is the formal uncertainty of that observation, derived on the basis of the signal-to-noise ratio achieved from fringe-fitting and ionosphere calibration process (if applicable), and  $\sigma_s$  is a baseline-dependent additive noise calculated for each session. This additive noise was determined through an iterative procedure such that the reduced chi-squared<sup>4</sup> of the post-fit residuals on each baseline for each session is about unity. This scheme ensures that the reduced chi-squared is close to unity over each session and furthermore over the entire data set as well.

### 3.4. Analysis software

An important element of the preparatory work for ICRF3 was that the data analysis was accomplished with several independent VLBI software packages running concurrently, one of which was also run separately at three institutions. This allowed the working group to check the results derived from these software packages against each other, which was essential to expose any issues, and to gain confidence in the overall data analysis scheme while the work progressed, in particular through the generation of several successive ICRF3 prototypes. Those software packages are the following: CALC-SOLVE (Caprette et al. 1990, see appendix), MODEST (Sovers & Jacobs 1996), OCCAM (Titov et al. 2004), QUASAR (Kurdubov 2007), VieVS (Böhm et al. 2018), and VieVS@GFZ (Nilsson et al. 2015). A detailed description of these software packages is beyond the scope of this paper. In all, they use very similar modeling (as described above). However, they are based on different estimation methods. CALC-SOLVE and VieVS use classical least-squares, while the other software packages employ specialized least-squares or filtering methods – MODEST is based on the square-root-information filter, OCCAM and QUASAR on the least squares collocation technique, and VieVS@GFZ on Kalman filtering.

<sup>4</sup> The reduced chi-squared  $\chi_\nu^2$  is defined as  $\chi^2/\nu$  where  $\chi^2$  is the sum of the squares of the weighted differences between the observed and calculated quantities and  $\nu$  is the degree of freedom which equals the number of observations minus the number of parameters.



While having such different software packages in hand was of importance for the ICRF3 preparatory work, including numerous tests, the final ICRF3 product was derived from data processing with a unique software package at each frequency band. The *S/X* band and *K* band data sets were analyzed with CALC-SOLVE while the *X/Ka* band data set was analyzed with MODEST. Combination of results from different software packages, for example by combining the individual normal equations, was investigated as an alternate option. However, it requires fine tuning of the input normal equations so that the estimated parameters are defined in a truly identical way in each software package and have consistent a priori settings, which was not possible to achieve within the available time to produce ICRF3 due to software limitations and other issues. We thus decided to not follow that option and instead to prefer individual determinations.

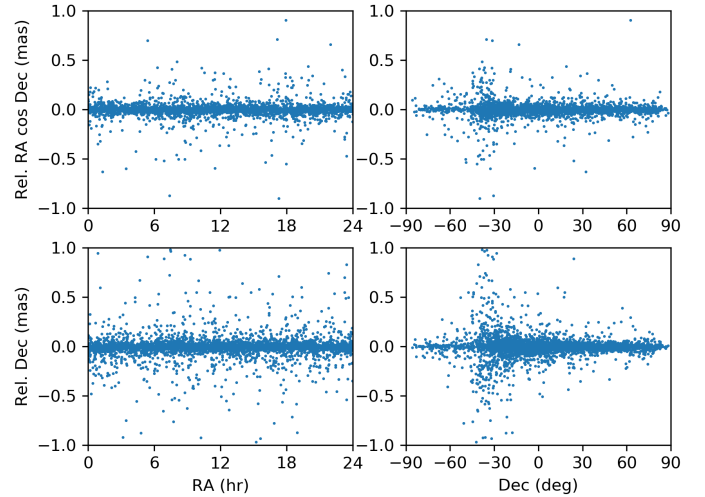
The reason why the *X/Ka* band data set was processed with a different software package than the *S/X* band and *K* band data sets was mainly a matter of convenience due to the specificity of the *X/Ka* band observations which are acquired, correlated, and post-processed through an entirely separate path. The two software packages used in the final data analysis for ICRF3, CALC-SOLVE and MODEST, have been inter-compared, although not recently, and were found to agree at the 1 ps level, as reported in Ma et al. (1998). Therefore, we are confident that proceeding this way does not limit the consistency of the ICRF3 results.

## 4. Assessment of errors

### 4.1. Variations in modeling and analysis configuration

An important part of the analysis work consisted in assessing the impact of some of the key features adopted for ICRF3 in terms of data selection, modeling or analysis configuration. To this end, a number of alternate analysis solutions were carried out, each of which varying a feature of interest. These alternate solutions were then checked against our original analysis solution, in particular by examining changes in the source positions, which allowed us to quantify the impact of those features on our results, including the level of systematic errors. All such tests were accomplished based on the *S/X* band data set and employed the CALC-SOLVE software package to guarantee the consistency of those alternate solutions with our original analysis solution.

An initial element that was investigated is the elevation cutoff angle for the data to be included into the analysis solutions. Observing at low elevations is necessary to decorrelate estimates of station vertical position from estimates of tropospheric zenith and gradient delays which both depend on the sine of the observation elevation. On the other hand, troposphere modeling errors increase with decreasing elevation because the observed signal passes through increasingly more troposphere. Below some elevation, the expected geometric improvement is overcome by modeling errors, which can thus bias the parameter determination (see, e.g., Herring 1986; Davis et al. 1991; MacMillan & Ma 1994). Figure 4 compares the estimated radio source coordinates when changing the cutoff elevation angle from 5° to 7°. As a matter of convenience, the resulting variations (and those for the additional tests presented below) are reported as  $\Delta\alpha \cos \delta$  and  $\Delta\delta$ . While the plots in Fig. 4 show differences up to 1 mas, only a minority of sources are subject to such large differences. The rms scatter is much lower, 6  $\mu$ as in  $\Delta\alpha \cos \delta$  and 8  $\mu$ as in  $\Delta\delta$ , with reduced chi-squared values of 0.06 and 0.07, respectively, hence indicating that the position variations are statistically not significant. Analysis solutions using



**Fig. 4.** Variations in the estimated radio source coordinates at *S/X* band when the observation elevation cutoff angle is changed from 5° to 7°. The differences are given as  $\Delta\alpha \cos \delta$  (upper panels) and  $\Delta\delta$  (lower panels) and are plotted as a function of right ascension (left-hand panels) and declination (right-hand panels). Units are milliarcseconds.

even higher cutoff elevation angles (10° and 15°) were also carried out to further characterize the impact. These led to rms scatters up to 20  $\mu$ as with reduced chi-squared up to 0.4, still below the expected coordinate uncertainties. Results of these tests are reported in Table 2.

The other elements of the modeling and analysis configuration that were investigated as part of this assessment include (i) the parametrization of the troposphere, approached by changing the interval between successive estimation of the zenith tropospheric delays from 30 min (the adopted interval) to shorter (20 min) or longer (1 hour and 3 hours) intervals, (ii) the treatment of the ICRF2 special handling sources, which was tackled by estimating the positions of these sources separately for each session (as for ICRF2) instead of uniquely from the entire data set, and (iii) the estimation of session-based antenna positions instead of global positions and velocities from all data. The results of these alternate analysis solutions are reported in Table 2 in terms of source coordinate variations and reduced chi-squared values in the same way as those regarding the cutoff elevation angle discussed above. In all, it is found that the source coordinate changes do not exceed 20  $\mu$ as, while reduced chi-squared values remain lower than 0.4. Based on these tests, it therefore appears that the corresponding choices of parametrization are not a major source of error in the realization of the frame.

A final element of the analysis configuration that was tested is the choice of the set of ICRF2 defining sources to be included in the no-net-rotation constraint applied for the alignment of the *S/X* band frame onto ICRF2. The reason for this test is that a fraction of the ICRF2 defining sources (about 10%) were found to have extended structures in post-ICRF2 VLBI imaging work. These sources either had not been imaged at the time ICRF2 was built or were subject to structural evolution in the meantime. A notable case is the source 0805+406 which shows a double structure with a component separation of about 6 mas, as revealed by VLBI images from the Bordeaux VLBI Image Database (BVID)<sup>5</sup> over the period 2010–2016. Its

<sup>5</sup> Accessible online at <http://bvid.astrophys.u-bordeaux.fr>.

**Table 2.** Impact of alternate analysis configurations on the estimated radio source coordinates at S/X band.

Elements of variations	$\Delta\alpha \cos \delta$		$\Delta\delta$	
	rms ( $\mu\text{as}$ )	$\chi^2_\nu$	rms ( $\mu\text{as}$ )	$\chi^2_\nu$
Elevation angle cutoff <sup>(a)</sup>				
>7°	6	0.06	8	0.07
>10°	10	0.18	14	0.20
>15°	16	0.36	22	0.38
Zenith tropospheric delays <sup>(b)</sup>				
20 min intervals	10	0.23	12	0.21
1 hr intervals	11	0.15	15	0.15
3 hr intervals	11	0.15	15	0.15
Special handling sources coordinates per session	13	0.38	9	0.13
Session-based antenna positions	8	0.11	11	0.14

**Notes.** The observed variations are characterized by the rms scatter (in microarcseconds) and reduced chi-squared for right ascension ( $\Delta\alpha \cos \delta$ ) and declination ( $\Delta\delta$ ). <sup>(a)</sup>The reference setting for the elevation angle cutoff is 5°. <sup>(b)</sup>The reference setting for the tropospheric delay intervals is 30 min.

X band structure index<sup>6</sup> in BVID has a median value of 5.5, well above the upper limit of 3.0 adopted for the selection of the ICRF2 defining sources. Such a source has very poor astrometric suitability and is improper as a defining source. Additionally, another six ICRF2 defining sources which have a structure index in BVID between 3.5 and 4.0 (0440+345, 1038+528, 1548+056, 1823+689, 2106–413, and 2326–477) were deemed to be not good enough either. In all, if considering all ICRF2 defining sources, 36 of them would not qualify anymore as defining sources according to the original ICRF2 criterion. It should be noted though that the bulk of these (24 sources) come with a structure index value between 3.0 and 3.25, which is just above the upper limit of 3.0 adopted for this selection. In order to test the potential degradation of the orientation of the frame due to these structured sources, we ran four alternate analysis solutions, each leaving out an increased number of ICRF2 defining sources from the constraint, depending on the maximum structure index considered as acceptable (4.0, 3.5, 3.25, 3.0). Since the data, modeling and estimated parameters were the same in the four solutions, the resulting frames then should differ only in orientation due to the slightly different no-net-rotation constraints applied. These four alternate frames were then compared to the frame obtained when all 295 ICRF2 defining sources are included in the constraint. We used for this purpose the formalism developed in Sect. 6.1 below. As reported in Table 3, the components of the rotations along the three axes,  $R_1$ ,  $R_2$ , and  $R_3$ , as derived from those comparisons, reach 6  $\mu\text{as}$  at most, which is below the estimated 10  $\mu\text{as}$  directional stability of ICRF2 (Fey et al. 2015) that defines the level to which the two frames should be aligned to be consistent in orientation. We thus concluded that there is no need to worry about these sources, hence our decision to keep all ICRF2 defining sources in the alignment process of ICRF3 onto ICRF2.

**Table 3.** Variations in the orientation of the S/X band frame depending on the subset of ICRF2 defining sources used to align the frame onto ICRF2.

Max SI	ICRF2 defining sources			Rotation of the frame		
	Nb excluded	Nb included	% included	$R_1$ ( $\mu\text{as}$ )	$R_2$ ( $\mu\text{as}$ )	$R_3$ ( $\mu\text{as}$ )
5.5	0	295	100	–	–	–
4.0	1	294	>99	–1	–1	4
3.5	7	288	98	–3	–2	4
3.25	12	283	96	–2	–2	6
3.0	36	259	88	–4	4	4

**Notes.** Each subset results from filtering out the ICRF2 defining sources based on the maximum structure index (SI) value given in column 1. Rotations (in  $\mu\text{as}$ ) are measured relative to the case where all 295 ICRF2 defining sources are considered for the alignment.

#### 4.2. Accuracy of estimated auxiliary parameters

An indirect way to assess the quality of our analysis solutions is to control the accuracy of the auxiliary parameters that were estimated simultaneously with the source coordinates and are thus an integral part of the solutions. The EOP are of particular value in this regard since estimates of these parameters are available for every session, which in some way allows one to check the entire analysis on a statistical basis. In order to evaluate the accuracy of the S/X band and K band EOP derived from our analysis solutions, we have compared them with those from the IVS combined series ivs15q2X<sup>7</sup>. The latter is an official product of the IVS obtained by combining individual EOP series generated by several IVS analysis centers. The comparison was achieved by differencing the EOP estimated for each session with those from the IVS combined series. The largest outliers ( $>5\sigma$ ) were then filtered out and a linear slope was fitted to the data and taken out to remove any long term trend between the series prior to computing statistics for the differences. This scheme follows the standard IERS practice for comparing and combining EOP series (see, e.g., Bizouard et al. 2019). It must be noted that only the post-1994 data were considered in this adjustment in order not to bias the assessment by the earlier less-accurate EOP determinations. After such processing, the weighted rms of the differences and reduced chi-squared were calculated as indicators of the agreement between our series and the ivs15q2X series. The results, including values of the fitted slopes, are given in Table 4 for the S/X band EOP series and in Table 5 for the K band EOP series. For completeness, median values of the formal uncertainties for the EOP, as derived from our solutions, and calculated over the same time span, are also reported in these tables.

By examining the content of Table 4, we first note that the median uncertainties for the five EOP are all about the same, in the range of 40–60  $\mu\text{as}$ , indicating that they are all determined with similar accuracies from the S/X band analysis solution. Looking at the results of the comparisons, we further note that the nutation offsets are in full agreement with those reported by the IVS (reduced chi-squared values of 1.0 for X and Y). The relative slopes estimated for these parameters are less than 1  $\mu\text{as yr}^{-1}$ , which is not significant at the reported accuracies of 50  $\mu\text{as}$ . The derived S/X band frame thus appears to be fully consistent with the currently available IVS nutation series. Differences for polar motion and UT1 are somewhat larger, with

<sup>6</sup> The structure index is an indicator of the astrometric suitability of the sources. See Fey & Charlot (1997) and Fey et al. (2015) for its definition and details on how it is calculated from VLBI source maps.

<sup>7</sup> As available online on August 25, 2018 from the IVS products web page at <https://ivscc.gsfc.nasa.gov/products-data/products.html>.

**Table 4.** Comparison of the EOP estimated at  $S/X$  band with those from the IVS combined series ivs15q2X.

Statistics for EOP (3464 data points)	$x_p$ ( $\mu\text{as}$ )	$y_p$ ( $\mu\text{as}$ )	UT1 ( $\mu\text{s}$ )	$X$ ( $\mu\text{as}$ )	$Y$ ( $\mu\text{as}$ )
Median uncertainty	61	56	2.6	55	56
Difference with IVS series					
wrms	76	79	6.1	45	44
$\chi^2_v$	2.2	2.4	2.6	1.0	1.0
slope <sup>(a)</sup>	6.1	10.3	-0.50	0.3	0.8
slope error <sup>(a)</sup>	$\pm 0.2$	$\pm 0.2$	$\pm 0.04$	$\pm 0.1$	$\pm 0.1$

**Notes.** Differences between our series and the IVS ones are characterized in terms of the weighted rms, reduced chi-squared and relative slope for each of the five EOP, i.e., the two polar motion components ( $x_p$ ,  $y_p$ ), the daily rotation UT1, and the two nutation offsets ( $X$ ,  $Y$ ). Median values of the EOP uncertainties resulting from our ICRF3  $S/X$  band analysis solution are also indicated in the table for comparison. <sup>(a)</sup>Units are  $\mu\text{as yr}^{-1}$  for ( $x_p$ ,  $y_p$ ),  $\mu\text{s yr}^{-1}$  for UT1, and  $\mu\text{as yr}^{-1}$  for ( $X$ ,  $Y$ ).

weighted rms of  $75\text{--}90\mu\text{as}$  and reduced chi-squared values in the range 2.2–2.6. Such increased differences may be due to a slight inconsistency (at the level of 1–2 mm) in fixing the origin and orientation of the terrestrial frame, which may slightly vary depending on the particular set of stations used in the constraints. Additionally, the relative slopes for these parameters are also found to have higher values, up to  $10\mu\text{as yr}^{-1}$  for polar motion component  $y_p$ . Again, this may be due to a slight inconsistency in fixing the rotation rate of the terrestrial frame. The consequence is a drift of the series on the long term, although at a limited level (7–12 mm shifts after 40 years). In order to further assess the matter, we have taken an additional step and compared our EOP series to the IERS EOP 14C04 series (Bizouard et al. 2019). The slopes derived for  $x_p$ ,  $y_p$ , and UT1 from this comparison (in units of  $\mu\text{as yr}^{-1}$  for  $x_p$  and  $y_p$ , and  $\mu\text{s yr}^{-1}$  for UT1) are 9.1, 9.4, and  $-1.64$ , to be checked against the results of the comparisons with the ivs15q2X series in Table 4, which are 6.1, 10.3, and  $-0.50$ . This indicates relative drifts of 3.0,  $-0.9$ , and  $-1.14$  between the IERS EOP 14C04 series and the ivs15q2X series. Calculating the mean of these drifts for the two components of polar motion and UT1, we derive a value of  $7\mu\text{as yr}^{-1}$ , whereas a value of  $8\mu\text{as yr}^{-1}$  is inferred when comparing our  $S/X$  band series with the ivs15q2X series. The inconsistency between our  $S/X$  band series and the IVS combined series, if any, is thus no larger than that between the IERS EOP 14C04 series and the IVS combined series. In any case, those small drifts have no impact on the derived celestial frame since they just reflect the slightly different ways in which the terrestrial frame was fixed and are expected to cancel out when the terrestrial frame and EOP relative rotations are added.

The results of comparing the EOP derived at  $K$  band with those from the ivs15q2X series, which are reported in Table 5, are based on a much smaller number of data points (less than 1% of the number of data points available at  $S/X$  band), possibly not statistically meaningful, and thus should be treated with caution. Nevertheless, such results are useful in providing indications on the quality and level of agreement of the estimated EOP at  $K$  band. As indicated by Table 5, median uncertainties are in the range between 40 and  $80\mu\text{as}$ , which is roughly at the same level as the median uncertainties at  $S/X$  band (see Table 4), with the exception of that found for parameter  $y_p$  which is much larger ( $140\mu\text{as}$ ). The latter likely originates in the geometrical configuration of the observing network at  $K$  band (i.e.,

**Table 5.** Comparison of the EOP estimated at  $K$  band with those from the IVS combined series ivs15q2X.

Statistics for EOP (27 data points)	$x_p$ ( $\mu\text{as}$ )	$y_p$ ( $\mu\text{as}$ )	UT1 ( $\mu\text{s}$ )	$X$ ( $\mu\text{as}$ )	$Y$ ( $\mu\text{as}$ )
Median uncertainty	66	140	5.5	43	42
Difference with IVS series					
wrms	153	246	14.9	88	101
$\chi^2_v$	4.7	6.5	10.2	3.7	5.5
slope <sup>(a)</sup>	15.6	20.1	0.51	$-1.0$	$-6.8$
slope error <sup>(a)</sup>	$\pm 8.1$	$\pm 11.4$	$\pm 1.65$	$\pm 3.5$	$\pm 3.9$

**Notes.** Differences between our series and the IVS ones are characterized in terms of the weighted rms, reduced chi-squared and relative slope for each of the five EOP, i.e., the two polar motion components ( $x_p$ ,  $y_p$ ), the daily rotation UT1, and the two nutation offsets ( $X$ ,  $Y$ ). Median values of the EOP uncertainties resulting from our ICRF3  $K$  band analysis solution are also indicated in the table for comparison. <sup>(a)</sup>Units are  $\mu\text{as yr}^{-1}$  for ( $x_p$ ,  $y_p$ ),  $\mu\text{s yr}^{-1}$  for UT1, and  $\mu\text{as yr}^{-1}$  for ( $X$ ,  $Y$ ).

the VLBA), which does not favor the estimation of that parameter. Comparisons with the ivs15q2X series show differences of  $100\text{--}250\mu\text{as}$ , which is twice as large as the differences found at  $S/X$  band, reaching a factor of three for parameter  $y_p$ . Such differences also appear to be fairly significant, with reduced chi-squared values of 4–10, reflecting the presence of systematic errors. In this regard, it is to be pointed out that our  $K$  band EOP estimates are fully independent of those in the IVS series since they are derived from different data sets, which is not the case of our  $S/X$  band estimates. For the same reason, interpolation errors may also be larger because there may not be an  $S/X$  band session carried out concurrently with each  $K$  band session. As regards polar motion and UT1, differences may also stem from the fixing of the origin and orientation of the terrestrial frame, which may not be attached to ITRF2014 as precisely as at  $S/X$  band, due to the use of only nine stations, all located in the Northern Hemisphere, for this purpose (see Sect. 3.3). On the other hand, the fitted slopes, even if they show higher values than at  $S/X$  band, do not have a significance that exceeds  $2\sigma$ . We have thus no indication of significant drifts with respect to ITRF2014. As noted previously, such drifts, in any case, would be absorbed and would have no impact on the derived celestial frame.

#### 4.3. Determination of realistic uncertainties

The formal uncertainties that come from geodetic and astrometric VLBI analyses are dependent on a number of factors, including the sensitivity of the network, the source flux density, and the number of observations. They get smaller as the sensitivity of the instrumentation (receivers, data acquisition terminals) improves or when using larger antennas, while they deteriorate as the sources become weaker (see, e.g., Malkin 2016, for an investigation of the relationship between source position uncertainty and flux density). Furthermore, the use of the least-squares method to solve for parameters in these analyses implies that the resulting formal uncertainties fall off as the square root of the number of observations. Consequently, such formal uncertainties may become very small, and even reach unrealistic levels, when the number of observations gets large. The number of observations in itself scales with the square of the number of antennas in the array. Larger VLBI arrays are thus more likely to generate formal uncertainties that are too small. Additionally, the proportion of independent VLBI measurements decreases as the size of



the array increases. For example, a five-station array (providing ten baselines) delivers ten VLBI delay measurements for a given source scan, among which only five, or 50%, are fully independent, a percentage that falls down to 22% for a ten-station array. Not accounting for the induced correlations (in particular related to clocks and troposphere) in the least-squares analysis then results in formal parameter uncertainties that are smaller than they should be in reality considering those correlations. The question of the validity of the formal uncertainties coming out from VLBI analyses has been debated ever since the beginning of VLBI. It was first investigated by [Ryan et al. \(1993\)](#) who concluded that a scaling factor of 1.5 is to be applied to VLBI formal uncertainties to bring them closer to actual uncertainties. Based on this finding, [Ma et al. \(1998\)](#) and [Fey et al. \(2015\)](#) also inflated the formal uncertainties from their analysis solutions by a factor of 1.5 to report ICRF and ICRF2 source coordinate uncertainties. We have not attempted here to redetermine this scaling factor and went along the same lines, in the continuity of the previous ICRF realizations. The scaling factor of 1.5, however, was only applied to the *S/X* band and *K* band source coordinate uncertainties, the *X/Ka* coordinate uncertainties being less likely to be affected by such underestimation because the *X/Ka* network consists mostly of single-baseline sessions, as indicated above.

Apart from this scaling factor, a noise floor was also applied to the source coordinate uncertainties reported in ICRF and ICRF2 so that these do not drop to unrealistic levels when the number of observations for a given source becomes very large. This noise floor was 250  $\mu\text{as}$  for ICRF and 40  $\mu\text{as}$  for ICRF2. In practice, those values were added in quadrature to the estimated formal uncertainties, scaled as described above, to derive the final source coordinate uncertainties. The determination of the noise floor in ICRF and ICRF2 was built upon comparisons of catalogs produced from independent analyses and/or independent data sets. In particular, [Fey et al. \(2015\)](#) ran a decimation test in which the ICRF2 data set was divided into two independent subsets and compared the source coordinates estimated from each subset of data to infer the noise level. We have followed a similar approach for ICRF3. To this end, all VLBI sessions were ordered temporally and divided into two subsets selected by even and odd sessions. The source coordinates estimated from the data in each subset were then compared in declination bands of 10°, limiting the comparison to sources with a minimum of 100 observations and observed in at least two sessions. From this comparison, the noise floor in each declination band was derived as the weighted rms of the coordinate differences for the sources in the given declination band divided by the square root of 2. That calculation was achieved separately for right ascension and declination. The reason for treating declination bands separately was to check the overall consistency of the process and to make sure no declination bands stick out, which may occur due to the non-uniform distribution of sources in declination (see Sect. 5.1 below). Such a scheme was applied to both the *S/X* band and *K* band data sets but not to the *X/Ka* band data set, considering that the amount of data available for the latter was not sufficient to make it meaningful. From our experience in processing data at *X/Ka* band, we decided instead to base the noise floor at this frequency band on that at *S/X* band.

The results of the decimation tests at *S/X* band show that the noise floor in the various declination bands ranges from 13 to 45  $\mu\text{as}$  in right ascension and from 12 to 64  $\mu\text{as}$  in declination (with four declination bands comprising less than 20 sources excluded from that assessment). Taking all sources together (i.e., without separating sources into declination bands), the above

**Table 6.** Scaling factor and noise floor (in  $\mu\text{as}$ ) applied to the estimated formal uncertainties of the source coordinates at each frequency band.

Frequency band	Right ascension		Declination	
	Scaling factor	Noise floor ( $\mu\text{as}$ )	Scaling factor	Noise floor ( $\mu\text{as}$ )
<i>S/X</i>	1.5	30	1.5	30
<i>K</i>	1.5	30	1.5	50
<i>X/Ka</i>	1.0	30	1.0	30

**Notes.** Inflated coordinate uncertainties are derived with the expressions  $(\sigma_\alpha \cos \delta)^2 = (s_\alpha \sigma_{\alpha,\text{formal}} \cos \delta)^2 + \sigma_{\alpha \cos \delta,0}^2$  and  $\sigma_\delta^2 = (s_\delta \sigma_{\delta,\text{formal}})^2 + \sigma_{\delta,0}^2$ , where  $\sigma_{\alpha,\text{formal}}$  and  $\sigma_{\delta,\text{formal}}$  are the formal uncertainties in right ascension and declination (i.e., resulting from the least-squares parameter adjustment in our VLBI analysis solutions) and  $\sigma_\alpha$  and  $\sigma_\delta$  are the corresponding inflated uncertainties. The scaling factors are expressed as  $s_\alpha$  and  $s_\delta$ , and the noise floor as  $\sigma_{\alpha \cos \delta,0}$  and  $\sigma_{\delta,0}$ .

numbers average to 26  $\mu\text{as}$  in right ascension and 30  $\mu\text{as}$  in declination. A question that arose was whether to treat differently the sources south of  $-40^\circ$  declination since there appears to be a disparity in the noise floor measured for these sources compared to those further north (44  $\mu\text{as}$  vs 25  $\mu\text{as}$  for right ascension and 47  $\mu\text{as}$  vs 29  $\mu\text{as}$  for declination). However, we did not do so considering that the assessment at those low declinations relies on far less sources (about 8% of the total number of usable sources from the decimation test) and that our estimation of the noise floor in the far south only differs by a factor of 1.7 from that further north. Based on these findings, we decided to adopt a noise floor of 30  $\mu\text{as}$  for both coordinates (right ascension and declination) without a declination dependency. Applying a similar scheme to the *K* band data led to values of the noise floor ranging from 18 to 57  $\mu\text{as}$  in right ascension and from 32 to 113  $\mu\text{as}$  in declination, with corresponding average values of 33  $\mu\text{as}$  and 57  $\mu\text{as}$ , respectively. Still, it must be noted that declinations below  $-30^\circ$  could not be assessed with our method because of too few data. Unlike at *S/X* band, the decimation tests at *K* band indicate that there is a notable difference in the noise floor between right ascension and declination. For this reason, we decided to adopt two different values, 30  $\mu\text{as}$  for right ascension and 50  $\mu\text{as}$  for declination. Table 6 summarizes the adopted values for the scaling factor and noise floor for the three frequency bands.

## 5. A multi-frequency frame

### 5.1. Three-frequency VLBI source positions

The modeling and analysis configuration outlined in Sect. 3 have been applied to the data sets at *S/X* band, *K* band, and *X/Ka* band described in Sect. 2 to produce three separate catalogs in a consistent way and all aligned onto the ICRS. These three catalogs form ICRF3, the first multi-frequency celestial reference frame ever realized. Some statistics about the data and fits performed at the three frequency bands, including the number of observations, their time span, the weighted rms of the post-fit residuals, and the reduced chi-squared are provided in Table 7, while the major features of ICRF3 (number of sources and median coordinate uncertainties at each frequency band) are summarized in Table 8. The *S/X* band catalog comprises 4536 sources, one-third more than in ICRF2, while the *K* band and *X/Ka* band catalogs comprise 824 sources and 678 sources, respectively. The diagram in Fig. 5 shows the breakdown of the sources according

**Table 7.** Overview of the data and least-squares fits accomplished at the  $S/X$ ,  $K$ , and  $X/Ka$  frequency bands, including statistics (weighted rms and reduced chi-squared) for the delay and delay rate residuals.

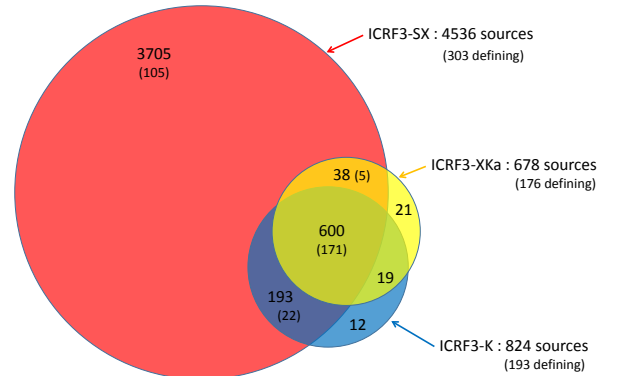
Frequency band	Number of observations	Delay residuals		Delay rate residuals		Data span
		wrms (ps)	$\chi^2_\nu$	wrms (fs s <sup>-1</sup> )	$\chi^2_\nu$	
$S/X$	13 190 274	25.5	1.05	–	–	1979 Aug 03 – 2018 Mar 27
$K$	482 616	17.5	0.91	66.6	0.90	2002 May 15 – 2018 May 05
$X/Ka$	69 062	43.6	1.00	102.3	1.00	2005 Jul 09 – 2018 Jan 28

**Table 8.** Major features of ICRF3, including the number of sources at each frequency band and statistics about coordinate uncertainties, correlation coefficients between right ascension and declination, and the error ellipse size (semi-major axis).

Frequency band	Number of sources	Statistics for all sources				Statistics for the common sources			
		Coordinate uncertainty $\alpha \cos \delta$ ( $\mu$ as)	Coordinate uncertainty $\delta$ ( $\mu$ as)	Correlation coeff.	Error ellipse ( $\mu$ as)	Coordinate uncertainty $\alpha \cos \delta$ ( $\mu$ as)	Coordinate uncertainty $\delta$ ( $\mu$ as)	Correlation coeff.	Error ellipse ( $\mu$ as)
$S/X$	4536	127	218	0.13	223	48	64	0.08	64
$K$	824	74	136	0.30	139	68	132	0.31	135
$X/Ka$	678	76	104	0.43	115	69	100	0.44	108

**Notes.** Statistics are provided for the entire  $S/X$ ,  $K$ , and  $X/Ka$  band catalogs and for the subset of 600 sources common to the three catalogs. All values reported in columns 3–10 are median values.

to frequency band. In all, ICRF3 includes a total 4588 sources, all of which are part of the  $S/X$  band catalog, except 52 sources which belong only to the  $K$  band and/or  $X/Ka$  band catalogs. Also to be noted is that 600 sources are common to the three catalogs. The median coordinate uncertainties in the  $S/X$  band catalog are 127  $\mu$ as for right ascension<sup>8</sup> and 218  $\mu$ as for declination, with the two coordinates generally only weakly correlated (median correlation coefficient of 0.13). The median of the error ellipse semi-major axis<sup>9</sup> is 223  $\mu$ as. These values represent a factor of 3.4 improvement compared to the ICRF2 figures where such median uncertainties were 396  $\mu$ as (for right ascension), 739  $\mu$ as (for declination), and 765  $\mu$ as (for the error ellipse semi-major axis). This dramatic improvement results from the VLBA campaigns that have been conducted since 2014 to re-observe all the less-observed ICRF2 sources (including the VCS sources), as pointed out in Sect. 2.1. Compared to the  $S/X$  band catalog, the median uncertainties for the  $K$  band and  $X/Ka$  band catalogs appear to be smaller by a factor of 1.5–2 (see the statistics for all sources in Table 8), while correlations between right ascension and declination appear to be stronger (median correlation coefficients of 0.30 and 0.43, respectively). The latter is likely a consequence of the more limited observing configurations at these two frequency bands. The observation that the median uncertainties are lower, however, is a somewhat artificial effect because the bulk of the sources in the  $S/X$  band catalog are VCS-type sources which in terms of positional precision are still not at the level of the most-observed (and most-precise)  $S/X$  band sources, despite the improvement noted above. When calculating median uncertainties solely for the 600 sources common to the three cat-

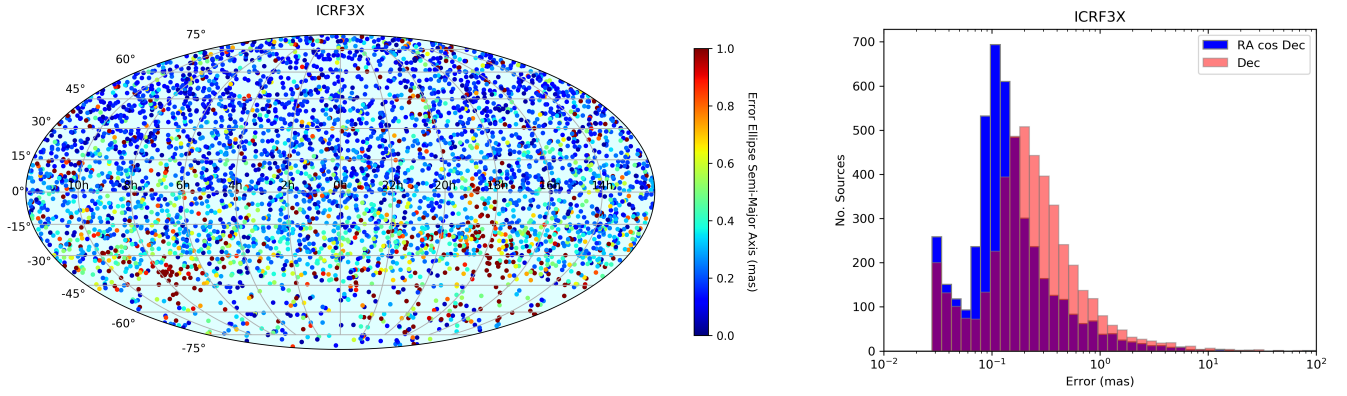
**Fig. 5.** Breakdown of the 4588 sources in ICRF3 according to frequency band. The circle colored red is for  $S/X$  band, the one colored blue is for  $K$  band, and the one colored yellow is for  $X/Ka$  band. The number of sources found in each colored area is printed within that area, with the number of ICRF3 defining sources (see Sect. 5.2) given in parentheses.

alogs (see the statistics for the common sources in Table 8), the figures are in fact the other way round. It is the  $S/X$  band coordinate uncertainties that turn out to be smaller by a factor of 1.5–2 compared to the  $K$  band and  $X/Ka$  band coordinate uncertainties. Notwithstanding this difference, the three catalogs show significantly improved positional precision compared to ICRF2.

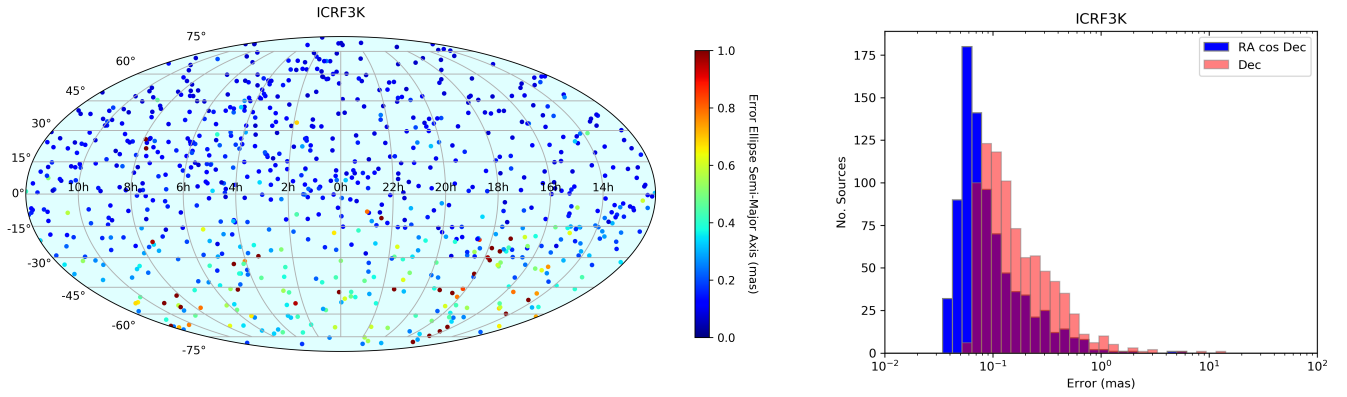
Figures 6–8 show the sky distribution and histogram of position uncertainties for the three catalogs. Not unexpectedly, the sky distribution for the  $S/X$  band and  $K$  band catalogs (left-hand panels in Figs. 6 and 7) has a deficiency of sources south of about  $-40^\circ$  declination (corresponding to the VLBA southern observing limit). The reason for this deficiency is that the VLBI stations able to observe further south are sparse. At the same time, the sky distribution for the  $X/Ka$  band catalog does not show such a division (see left-hand panel in Fig. 8). With only four sites, the VLBI network used at this frequency band is not tailored to

<sup>8</sup> Here and in the following sections, right ascension always denotes right ascension multiplied by the cosine of the declination.

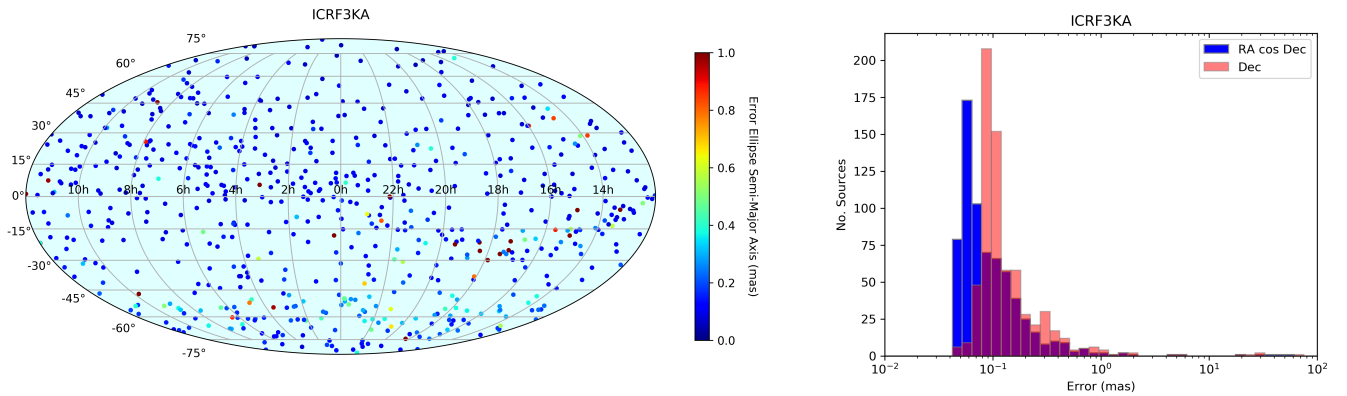
<sup>9</sup> The error ellipse semi-major axis is calculated from the right ascension and declination uncertainties and correlation coefficient between the coordinates. See, e.g., Eq. (1) of Gaia Collaboration (2018).



**Fig. 6.** *Left:* distribution of the 4536 sources included in the ICRF3 *S/X* band frame on a Mollweide projection of the celestial sphere. Each source is plotted as a dot color-coded according to its position uncertainty (defined as the semi-major axis of the error ellipse in position). *Right:* distribution of coordinate uncertainties for the same 4536 sources. Right ascension is shown in blue while declination is shown in salmon. The superimposed portion of the two distributions is shown in purple.



**Fig. 7.** *Left:* distribution of the 824 sources included in the ICRF3 *K* band frame on a Mollweide projection of the celestial sphere. Each source is plotted as a dot color-coded according to its position uncertainty (defined as the semi-major axis of the error ellipse in position). *Right:* distribution of coordinate uncertainties for the same 824 sources. Right ascension is shown in blue while declination is shown in salmon. The superimposed portion of the two distributions is shown in purple. It must be noted that the scale for the y-axis (number of sources) is different from that in Fig. 6.

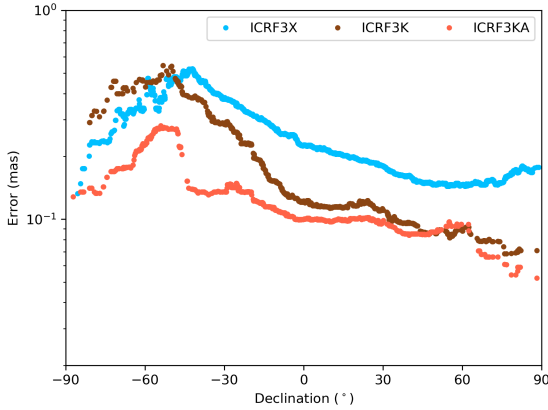


**Fig. 8.** *Left:* distribution of the 678 sources included in the ICRF3 *X/Ka* band frame on a Mollweide projection of the celestial sphere. Each source is plotted as a dot color-coded according to its position uncertainty (defined as the semi-major axis of the error ellipse in position). *Right:* distribution of coordinate uncertainties for the same 678 sources. Right ascension is shown in blue while declination is shown in salmon. The superimposed portion of the two distributions is shown in purple. It must be noted that the scale for the y-axis (number of sources) is different from that in Fig. 6.

densifying the frame (as opposed to the VLBA at the two other frequency bands), hence the more uniform (although not as dense) sky distribution at *X/Ka* band. The histograms of position uncertainties in the right-hand panels of Figs 6–8 show that the distributions of uncertainties for both the *K* band and

*X/Ka* band catalogs peak at about  $50\text{--}60\mu\text{s}$  in right ascension and  $80\text{--}90\mu\text{s}$  in declination, while at *S/X* band the peaks are near  $100\mu\text{s}$  for right ascension and  $200\mu\text{s}$  for declination. These numbers are in line with the median uncertainties reported in Table 8. The fact that the uncertainties in declination are a





**Fig. 9.** Median source position uncertainty as a function of declination for the *S/X* band frame (shown in light blue), the *K* band frame (shown in brown), and the *X/Ka* band frame (shown in salmon). The quantity represented is the running median for the semi-major axis of the error ellipse in position, determined over declination bins of  $15^\circ$  in each frame.

factor of 1.5–2.0 larger than those in right ascension finds its origin in the geometry of the VLBI observing networks which have longer east-west than north-south baselines (see Sect. 2). In Fig. 6, the histogram of position uncertainties also reveals that the *S/X* band distribution of uncertainties has a secondary peak. This peak is just above  $30 \mu\text{as}$ , at the noise floor, and captures the block of roughly 500 sources that have the most precise positions. It must be noted that no sources with VLBA-only observations are present in this block, meaning that they all have been observed as part of IVS programs (possibly jointly with the VLBA). In all, there are 503 sources in the pool of sources observed by the IVS that have a more precise position than any of the sources observed solely with the VLBA. For the latter, the position uncertainty (defined as the semi-major axis of the error ellipse in position) is  $64 \mu\text{as}$  at best and thus does not reach the noise floor. Looking at the other end of the distribution (i.e., where the least precise source positions are found), there are 359 sources with position uncertainty worse than 1 mas (corresponding to 8% of the *S/X* band catalog), including 22 sources (0.5% of the catalog) that have position uncertainties worse than 10 mas. These numbers contrast with those for ICRF2, where 1428 sources (42% of the catalog) had position uncertainties worse than 1 mas, 204 of which (6% of the catalog) were found to have position uncertainties worse than 10 mas. The *K* band and *X/Ka* band frames contain an even smaller portion of sources with less precise positions – 26 sources at *K* band (3% of the catalog) and 17 sources at *X/Ka* band (2.5% of the catalog) have position uncertainties worse than 1 mas, while only a handful of these have position uncertainties worse than 10 mas (one source at *K* band and five sources at *X/Ka* band). Based on the color coding of the position uncertainties in Figs. 6–8 (left-hand panels), one also sees that the distribution of uncertainties on the sky is not uniform. Overall, the southern sources have generally less precise positions than the northern ones. This is reflected in Fig. 9 which shows how the median position uncertainty varies as a function of declination in the three catalogs. While roughly stable at the highest declinations ( $> 40^\circ$ ), the *S/X* band median position uncertainty degrades regularly when declination goes from  $40^\circ$  to  $-45^\circ$ , after which it improves again toward  $-90^\circ$  declination. The *K* band and *X/Ka* band catalogs show similar properties, with median position uncertainty degrading from

$90^\circ$  to  $-50^\circ$  declination and improving further south, from  $-50^\circ$  to  $-90^\circ$  declination.

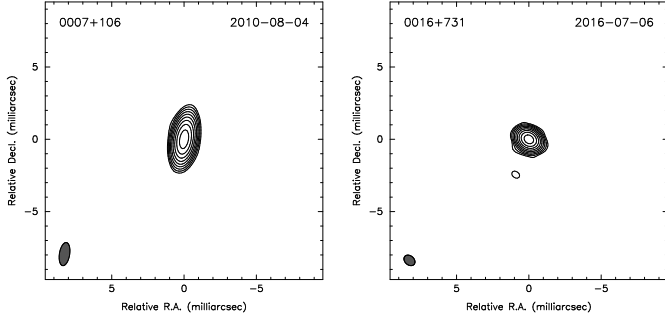
## 5.2. Selection of defining sources

Taking advantage of the much extended VLBI data set now available and acknowledging the fact that some of the ICRF2 defining sources were found to be no longer suitable as defining sources (see Sect. 4.1), we decided to select a new set of defining sources for ICRF3 based on specific criteria and not considering the set of ICRF2 defining sources as a starting set. For this purpose, three criteria were put forward: (i) the overall sky distribution of the ICRF3 defining sources, (ii) the position stability of the individual sources, and (iii) the compactness of their structures. While the second and third criteria were already considered for ICRF2, the first one (sky distribution) was not a major element of selection. The sole consideration in this regard consisted in splitting the celestial sphere into five declination bands and arranging for a similar number of defining sources to be selected in each band. This resulted in a distribution of the ICRF2 defining sources on the sky not fully optimum although covering the entire sky. Our goal this time was to work toward a set of ICRF3 defining sources that has a uniform distribution on the sky.

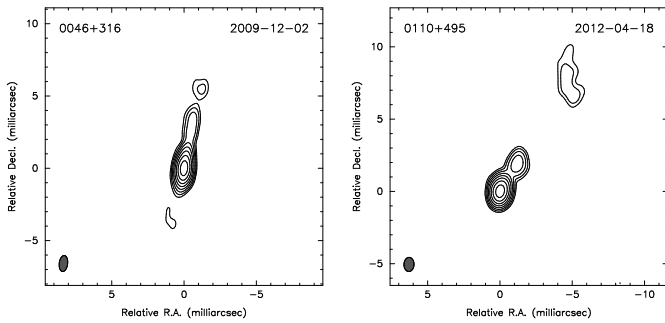
The scheme used to achieve a uniform source distribution consisted in dividing the celestial sphere into equivalent sectors and identifying the most suitable source in each sector based on criteria (ii) and (iii) above. The number of sectors was chosen as a compromise between having more defining sources and having a sufficient number of sources in each sector to identify at least one source with the required properties in terms of position stability and source structure. The choice was made for a total of 324 sectors, splitting right ascensions into 18 equal angular sections, each 1 h 20 min wide, while the *z*-axis, which joins the southern and northern poles (at  $-90^\circ$  and  $+90^\circ$  declination, respectively), was also cut into 18 equal segments. It must be noted that the latter implies that the separation of the declination bands increases toward the poles. All sources within each sector that had observations at *S/X* band in at least 20 sessions were then extracted and ranked according to positional stability. For this ranking, we followed the method employed for ICRF2 (see Fey et al. 2015) and defined the source position stability as

$$s = \sqrt{\text{wrms}_{\alpha \cos \delta}^2 \chi_{\nu, \alpha \cos \delta}^2 + \text{wrms}_{\delta}^2 \chi_{\nu, \delta}^2}, \quad (5)$$

where  $\text{wrms}_{\alpha \cos \delta}$  and  $\text{wrms}_{\delta}$  are the weighted rms of the coordinate variations about the weighted mean coordinates for right ascension and declination, respectively, and  $\chi_{\nu, \alpha \cos \delta}^2$  and  $\chi_{\nu, \delta}^2$  are the reduced chi-squared of the fit to the mean coordinates. The coordinate time series used in the calculation were obtained from four different analysis solutions, one where the positions of the ICRF2 defining sources were solved globally (i.e., estimated once from the entire data set), while the positions for the rest of the sources were estimated session-wise, and three others where the positions for one-third of the ICRF2 defining sources, selected in turn, were estimated session-wise, while the positions for all other sources were solved globally. In this scheme, no-net-rotation constraints were applied either to the full set of ICRF2 defining sources (in the first case) or only to two-thirds of it (corresponding to the portion of defining sources that was solved globally in the three other solutions), in order to fix the orientation of the frame. Following this ranking, the VLBI morphology of all extracted sources was closely examined to assess their compactness and hence suitability as defining sources. This was achieved through visual inspection of the multi-epoch BVID



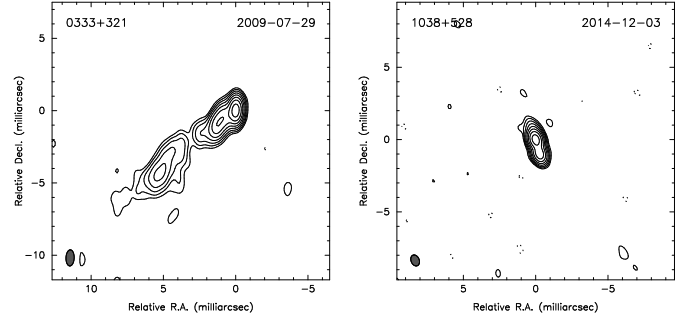
**Fig. 10.** VLBI maps at X band of two ICRF3 defining sources with structural morphology in category A. Contour levels are drawn at  $\pm 0.25$ , 0.5, 1, 2, 4, 8, 16, 32, and 64% of the peak brightness. The labels in each plot specify the source name and the epoch of the observations. These maps (retrieved from BVID) were made from data from RDV sessions.



**Fig. 11.** VLBI maps at X band of two ICRF3 defining sources with structural morphology in category B. Contour levels are drawn at  $\pm 0.5$ , 1, 2, 4, 8, 16, 32, and 64% of the peak brightness. The labels in each plot specify the source name and the epoch of the observations. These maps (retrieved from BVID) were made from data from RDV sessions.

images available, supplemented with those from the Radio Reference Frame Image Database<sup>10</sup>, looking in particular for structural variations, and by checking source structure indices and their variability with time (as available from BVID). Purpose-made VLBI images of sources in the Southern Hemisphere were also used to supplement the material from those two databases. Based on this assessment, the sources within each sector were separated into three categories, those that show minimal structure, usually qualified as point-like or quasi point-like (category A), those that show moderately extended structures (category B), and those that show extended or very extended structures (category C). From that categorization, the selection of the defining sources was then performed by choosing the top-ranked source from category A within each sector or in the case that there is no such source the top-ranked source from category B.

Following the above scheme, an initial pool of 702 sources having observations in at least 20 sessions in the S/X band catalog was identified as a subset of potential defining sources. These sources cover 322 sectors, thus leaving two sectors empty, while each of the other (non-empty) sectors includes between one and eight sources. Moving further with the source categorization, 216 sectors were found to contain at least one source from category A, while 62 sectors had only category B (and C) sources and 19 sectors had only category C sources. In the remaining



**Fig. 12.** VLBI maps at X band of two ICRF3 sources with structural morphology in category C. Contour levels are drawn at  $\pm 0.5$ , 1, 2, 4, 8, 16, 32, and 64% of the peak brightness. The labels in each plot specify the source name and the epoch of the observations. It should be emphasized that the source pictured in the right-hand panel has a double morphology with two closely-spaced components. These maps (retrieved from BVID) were made from data from RDV sessions.

25 sectors, source structure could not be assessed because of the lack of images. The choice was then made to leave out the 19 sectors with only category C sources since these have poor astrometric quality due to their having extended VLBI structures and cannot be deemed suitable as defining sources. On the other hand, the top-ranked sources in the 25 sectors where no images were available were kept in since there was no reason to assume that the structure of those sources would be inadequate. In all, this leaves a total of 303 defining sources where 216 of these (72%) have either good or excellent astrometric suitability (i.e., are category A sources) and 62 others (20%) have acceptable (if not ideal) quality (corresponding to category B sources), the astrometric suitability of the 25 remaining sources (8%) being unknown. See Table 18 below for the identification of the 21 sectors where no suitable defining source was found and Tables 17 and 19 for the identification of the defining sources for which structure was either not assessed or found to fall into category B. Figures 10 and 11 provide examples of images at X band for category A and category B defining sources. For comparison purposes, examples of images at X band of category C sources (excluded from the algorithm of selection for the defining sources) are shown in Fig. 12. Of interest is that among the 303 defining sources, 246 were ranked first (within their sector) in terms of source position stability, while 40 were ranked second, 12 were ranked third, four were ranked fourth, and one was ranked fifth, hence indicating good overall consistency between position stability and source compactness. In terms of observing characteristics, the vast majority of the selected sources fall into the pool of sources observed by the IVS – only 21 of them have observations that were conducted solely with the VLBA.

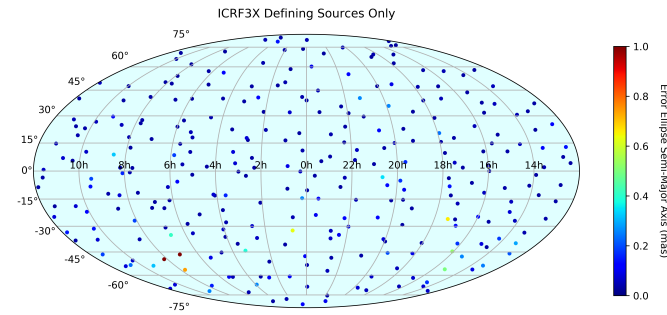
As expected, the sky distribution of the defining sources thus selected ends up being fairly uniform due to the scheme adopted for this selection (see Fig. 13). Looking at the distribution of coordinate uncertainties in Fig. 14, it is striking that the defining sources, in their vast majority, show very precise positions despite this quantity not being used as a selection criterion. It is also notable that the histograms of uncertainties in right ascension and declination are superimposed, a situation that differs from that observed when considering the entire S/X band catalog where the peaks for the uncertainties in right ascension and declination are shifted relative to one another (see Fig. 6). These specificities are further reflected by the values of the corresponding median uncertainties which are close to the noise floor and

<sup>10</sup> Accessible online at <https://www.usno.navy.mil/USNO/astrometry/vlbi-products/rrfid>.

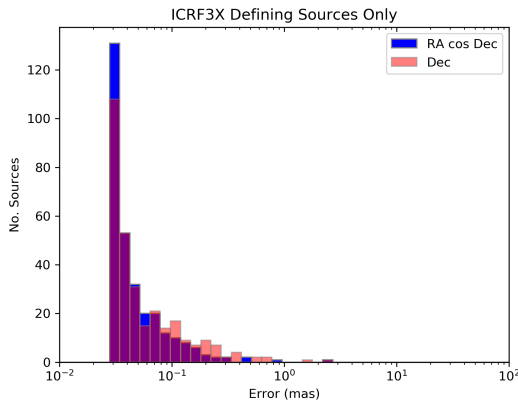
**Table 9.** Number of ICRF3 defining sources at each frequency band and statistics about coordinate uncertainties, correlation coefficients between right ascension and declination, and the error ellipse size (semi-major axis) for these sources.

Frequency band	Number of sources	Statistics for all defining sources				Statistics for the common defining sources			
		Coordinate uncertainty $\alpha \cos \delta$ ( $\mu\text{as}$ )	$\delta$ ( $\mu\text{as}$ )	Correlation coeff.	Error ellipse ( $\mu\text{as}$ )	Coordinate uncertainty $\alpha \cos \delta$ ( $\mu\text{as}$ )	$\delta$ ( $\mu\text{as}$ )	Correlation coeff.	Error ellipse ( $\mu\text{as}$ )
<i>S/X</i>	303	36	41	0.07	41	33	35	0.06	35
<i>K</i>	193	63	120	0.28	122	60	113	0.27	115
<i>X/Ka</i>	176	62	92	0.40	98	62	92	0.40	97

**Notes.** Statistics are provided for all defining sources in each catalog and for the 171 defining sources common to the three catalogs. All values reported in columns 3–10 are median values.



**Fig. 13.** Sky distribution of the 303 ICRF3 defining sources. Each source is plotted as a dot color-coded according to its position uncertainty in the *S/X* band frame (where the position uncertainty is defined as the semi-major axis of the error ellipse in position).



**Fig. 14.** Distribution of coordinate uncertainties of the 303 ICRF3 defining sources at *S/X* band. Right ascension is shown in blue while declination is shown in salmon. The superimposed portion of the two distributions is shown in purple.

similar for right ascension and declination (36  $\mu\text{as}$  and 41  $\mu\text{as}$ , respectively), as reported in Table 9. Correlation coefficients between right ascension and declination have a median value that is even smaller than when computed for the entire catalog (0.07 vs 0.13), therefore indicating very weak correlation between the two coordinates in general. At *K* band and *X/Ka* band, there are no such striking differences between the defining sources and the rest of the catalog sources. While median uncertainties appear to be slightly better for the defining sources (by about 10–15%), the magnitude of these uncertainties remains worse by a factor of two to three when compared to the *S/X* band median uncertainties (see Table 9). Median correlation coefficients do not show any significant differences either, whether computed for all cata-

log sources (Table 8) or solely for the subset of defining sources (Table 9). The absence of apparent specificities for the defining sources at *K* band or *X/Ka* band is not unexpected since the algorithm for selecting the defining sources was only tailored to the properties of the sources and configuration of the observations at *S/X* band. This finding may also reflect the fact that the sources tend to be more compact at higher frequencies (Charlot et al. 2010), thus mitigating the differences observed at *S/X* band (in terms of structural properties) between the defining and non-defining sources, or simply biases in the selection of the sources targeted for observation at *K* band and *X/Ka* band.

### 5.3. Practical use of the frame

The coordinates of the 4588 sources comprised in ICRF3 along with their uncertainties are given in Tables 10–12. Table 10 is for the *S/X* band frame, Table 11 is for the *K* band frame, and Table 12 is for the *X/Ka* band frame. Besides source coordinates, the three tables also include proper information to identify each source (ICRF designation and IERS name) and details about the VLBI sessions (first and last sessions in which a source was observed, mean epoch of the sessions, number of sessions), the observations (number of VLBI delays and delay rates used to estimate the source position), and the characteristics of the errors (correlation coefficient between right ascension and declination). The ICRF3 defining sources are also identified in these tables, as are those sources that were observed solely with the VLBA. The total number of sources in this condition (i.e., with VLBA-only data) is 3084 for the *S/X* band frame and 544 for the *K* band frame. For the sources that have been imaged, indicators about their astrometric suitability (i.e., source structure indices and compactness) are available from the BVID database. As supplemental information, the basic optical characteristics of most ICRF3 sources may be found in the Optical Characteristics of Astrometric Radio Sources catalog (Malkin 2018).

The ICRF3 source coordinates reported in Tables 10–12 are provided for epoch 2015.0. As explained above, these coordinates should be propagated for observations at epochs away from that epoch using a Galactic acceleration amplitude of 5.8  $\mu\text{as yr}^{-1}$ . In practice, this may be accomplished using the formulas

$$\alpha_t = \alpha_{t_0} + \Delta\mu_\alpha (t - t_0), \quad (6)$$

$$\delta_t = \delta_{t_0} + \Delta\mu_\delta (t - t_0), \quad (7)$$

where  $(\alpha_{t_0}, \delta_{t_0})$  are the ICRF3 source coordinates (i.e., the coordinates at epoch  $t_0 = 2015.0$ ), while  $(\alpha_t, \delta_t)$  are the coordinates at epoch  $t$ . The expressions for the components of the proper motion induced by Galactic acceleration ( $\Delta\mu_\alpha, \Delta\mu_\delta$ ) may



**Table 10.** Coordinates at epoch 2015.0 for the 4536 sources included in the S/X band frame.

Source identification		Cat. <sup>(c)</sup>	Right ascension (h m s)	Declination (° ' ")	Coordinate uncertainty		Correl. <sup>(d)</sup>	Epoch of sessions <sup>(e)</sup>			Observations <sup>(f)</sup>	
ICRF designation <sup>(a)</sup>	IERS name <sup>(b)</sup>				(s)	(")		Mean	First	Last	N <sub>ses</sub>	N <sub>del</sub> N <sub>rat</sub>
ICRF J000020.3–322101	2357–326	V	00 00 20.39997606	–32 21 01.2337415	0.00000804	0.0002624	–0.0602	56559.8	52306.7	57776.0	4	237 0
ICRF J000027.0+030715	2357+028		00 00 27.02251377	03 07 15.6463606	0.00005931	0.0003421	–0.0119	57974.7	57974.7	57974.7	1	28 0
ICRF J000053.0+405401	2358+406	V	00 00 53.08106320	40 54 01.8096518	0.00001504	0.0002670	–0.1654	56460.2	50242.8	57809.9	3	152 0
ICRF J000105.3–155107	2358–161	V	00 01 05.32873479	–15 51 07.0752302	0.00000702	0.0002261	–0.2106	56338.4	50632.3	58137.6	4	316 0
ICRF J000107.0+605122	2358+605	V	00 01 07.09981547	60 51 22.7980875	0.00003378	0.0001948	0.1619	57160.2	52306.7	57836.8	3	172 0
ICRF J000108.6+191433	2358+189		00 01 08.62156616	19 14 33.8017136	0.00000260	0.0000472	–0.0314	55771.9	50085.5	58205.8	168	3584 0

**Notes.** The coordinates reported in the table may be propagated at other epochs using Eqs. (6) and (7). The content printed here corresponds to the entries for the first six sources. The table in its entirety is available in electronic form at the CDS. <sup>(a)</sup>The ICRF designations were derived from the source coordinates with the format ICRF JHHMMSS.s+DDMMSS or ICRF JHHMMSS.s–DDMMSS, according to the original ICRF prescriptions. <sup>(b)</sup>The complete format for the IERS names, constructed from previous B1950.0 coordinates, includes acronym and epoch and is IERS BHHMM+DDd or IERS BHHMM–DDd. <sup>(c)</sup>The ICRF3 defining sources are identified with a “D” in this column. Additionally, the sources observed solely with the VLBA are marked with a “V”. <sup>(d)</sup>The value given in this column indicates the correlation coefficient between the estimated right ascension and declination coordinates. <sup>(e)</sup>The values given in these three columns are expressed as Modified Julian Date (MJD), i.e., JD – 2 400 000.5. <sup>(f)</sup>The three columns indicate the number of sessions in which the source was observed (N<sub>ses</sub>) and the number of delay and delay rate observations used to estimate its coordinates (N<sub>del</sub> and N<sub>rat</sub>).

**Table 11.** Coordinates at epoch 2015.0 for the 824 sources included in the K band frame.

Source identification		Cat. <sup>(c)</sup>	Right ascension (h m s)	Declination (° ' ")	Coordinate uncertainty		Correl. <sup>(d)</sup>	Epoch of sessions <sup>(e)</sup>			Observations <sup>(f)</sup>	
ICRF designation <sup>(a)</sup>	IERS name <sup>(b)</sup>				(s)	(")		Mean	First	Last	N <sub>ses</sub>	N <sub>del</sub> N <sub>rat</sub>
ICRF J000435.6–473619	0002–478	D	00 04 35.65553663	–47 36 19.6034797	0.00004466	0.0010827	0.6589	57961.2	57599.3	58067.8	7	13 13
ICRF J000504.3+542824	0002+541	V	00 05 04.36334104	54 28 24.9244365	0.00000840	0.0000905	0.0477	57561.2	53898.6	58181.5	9	732 732
ICRF J000557.1+382015	0003+380	V	00 05 57.17539180	38 20 15.1489842	0.00000450	0.0000737	–0.1988	57909.1	57372.1	58195.8	10	724 724
ICRF J000613.8–062335	0003–066		00 06 13.89288203	–06 23 35.3357441	0.00000321	0.0001061	–0.4903	57819.0	56782.4	58181.5	9	658 658
ICRF J000903.9+062821	0006+061	V	00 09 03.93185280	06 28 21.2397824	0.00000772	0.0002257	–0.3528	57837.4	52782.5	58195.8	7	316 316
ICRF J001031.0+105829	0007+106	DV	00 10 31.00590225	10 58 29.5043805	0.00000325	0.0000920	–0.3330	55132.2	52782.5	58244.2	12	1043 1043

**Notes.** The coordinates reported in the table may be propagated at other epochs using Eqs. (6) and (7). The content printed here corresponds to the entries for the first six sources. The table in its entirety is available in electronic form at the CDS. See Table 10 for the meaning of the footnotes.

**Table 12.** Coordinates at epoch 2015.0 for the 678 sources included in the X/Ka band frame.

Source identification		Cat. <sup>(c)</sup>	Right ascension (h m s)	Declination (° ' ")	Coordinate uncertainty		Correl. <sup>(d)</sup>	Epoch of sessions <sup>(e)</sup>			Observations <sup>(f)</sup>	
ICRF designation <sup>(a)</sup>	IERS name <sup>(b)</sup>				(s)	(")		Mean	First	Last	N <sub>ses</sub>	N <sub>del</sub> N <sub>rat</sub>
ICRF J000435.6–473619	0002–478	D	00 04 35.655546616	–47 36 19.6047565	0.00001340	0.0001584	–0.4333	57086.5	56297.6	58027.3	23	30 30
ICRF J000504.3+542824	0002+541		00 05 04.36338455	54 28 24.9241634	0.00001033	0.0001306	0.1878	57364.8	55304.5	58146.7	27	54 54
ICRF J000557.1+382015	0003+380		00 05 57.17541201	38 20 15.1487480	0.00000509	0.0000814	–0.3088	56844.3	53561.7	58146.9	88	126 126
ICRF J000613.8–062335	0003–066		00 06 13.89287585	–06 23 35.3351018	0.00000627	0.0001293	–0.5731	56545.8	53561.6	58146.8	92	137 137
ICRF J000903.9+062821	0006+061		00 09 03.93183737	06 28 21.2399247	0.00000619	0.0001190	–0.6048	56785.3	53651.5	58146.8	77	110 110
ICRF J001031.0+105829	0007+106	D	00 10 31.00590425	10 58 29.5042938	0.00000375	0.0000883	–0.4731	56132.3	53694.3	58146.9	113	179 179

**Notes.** The coordinates reported in the table may be propagated at other epochs using Eqs. (6) and (7). The content printed here corresponds to the entries for the first six sources. The table in its entirety is available in electronic form at the CDS. See Table 10 for the meaning of the footnotes.

be found, for instance, in [MacMillan et al. \(2019\)](#) and are given by

$$\Delta\mu_\alpha \cos \delta = -A_1 \sin \alpha + A_2 \cos \alpha, \quad (8)$$

$$\Delta\mu_\delta = -A_1 \cos \alpha \sin \delta - A_2 \sin \alpha \sin \delta + A_3 \cos \delta, \quad (9)$$

where the  $A_i$  parameters are the barycentric components of the Galactic acceleration vector, scaled by  $1/c$  as in Eq. (1). These may be expressed as  $A_G(\cos \delta_G \cos \alpha_G, \cos \delta_G \sin \alpha_G, \sin \delta_G)$ , where  $(\alpha_G, \delta_G)$  are the equatorial coordinates of the vector direction ( $\alpha_G = 266.4^\circ$ ,  $\delta_G = -29.0^\circ$ ) and  $A_G$  is the amplitude of this vector, as defined above ( $A_G = 5.8 \mu\text{as yr}^{-1}$ ). It is to be pointed out that propagation of the ICRF3 coordinates at other epochs using the above equations is required only for the most accurate needs. For observations within ten years of the reference epoch of the frame (2015.0), such adjustments may not be necessary unless a positional accuracy better than  $100 \mu\text{as}$  is desired.

Apart from such considerations on propagation of the source coordinates, the practical use of ICRF3 also implies that one has to select positions in one of the  $S/X$ ,  $K$ , or  $X/Ka$  band catalogs for sources that have positions available at more than one frequency band. In this respect, we recommend that the  $S/X$  band positions be used at first since these are on average more accurate than the  $K$  band or  $X/Ka$  band positions (see Tables 8 and 9). However, when positions at  $K$  band or  $X/Ka$  band are specifically desired, these should be favored over the  $S/X$  band positions. This is the case in particular when ICRF sources are used as calibrators to observe water masers at  $K$  band (e.g., [Immer et al. 2013](#)) or as fiducial references for spacecraft navigation at  $X/Ka$  band as deep space flights move to  $Ka$  band links ([Morabito 2017](#)).

## 6. Analysis of ICRF3

Aside from internal assessments, comparisons of ICRF3 with independent realizations of the extragalactic frame are essential to control its quality, its alignment onto the ICRS, and potential deformations of the frame. The two independent frames used for this purpose are the predecessor of ICRF3, ICRF2 ([Fey et al. 2015](#)), and the recently released *Gaia*-CRF2 frame in the optical domain ([Gaia Collaboration 2018](#)). Also important is the assessment of the consistency of the individual source positions at the three different ICRF3 frequencies and with optical ones in *Gaia*-CRF2 since the measured positions may depend at some level on the observed frequency band due to possible physical offsets of the emission in the different bands. All such checks are reported below, including a description of the approach devised for these comparisons, prior to the presentation of the results.

### 6.1. Modeling rotations and deformations between catalogs

The scheme that we used for comparing catalogs follows that developed by [Mignard & Klioner \(2012\)](#) and is based on vector spherical harmonics decomposition. In this scheme, the coordinate differences between catalogs are modeled by a transformation that takes the global rotation between the catalogs into account as well as the low-degree deformations. Mathematically, the corresponding coordinate transformation includes three rotations, three “glide” parameters that characterize the dipolar deformation of the coordinate field (see [Mignard & Klioner 2012](#)), and ten quadrupole terms representing degree-2 deformations. Adopting the notation of [Titov & Lambert \(2013\)](#), the coordinate differences between the catalogs may be expressed as

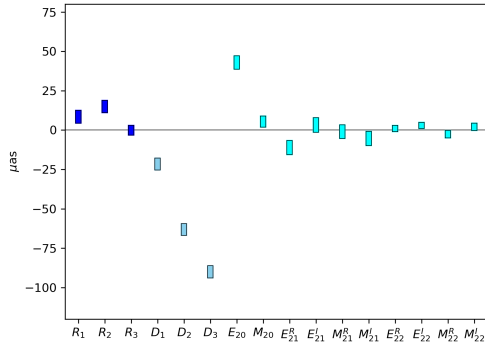
$$\begin{aligned} \Delta\alpha \cos \delta = & R_1 \cos \alpha \sin \delta + R_2 \sin \alpha \sin \delta - R_3 \cos \delta \\ & - D_1 \sin \alpha + D_2 \cos \alpha \\ & + M_{20} \sin 2\delta \\ & + (E_{21}^{\text{Re}} \sin \alpha + E_{21}^{\text{Im}} \cos \alpha) \sin \delta \\ & - (M_{21}^{\text{Re}} \cos \alpha - M_{21}^{\text{Im}} \sin \alpha) \cos 2\delta \\ & - 2(E_{22}^{\text{Re}} \sin 2\alpha + E_{22}^{\text{Im}} \cos 2\alpha) \cos \delta \\ & - (M_{22}^{\text{Re}} \cos 2\alpha - M_{22}^{\text{Im}} \sin 2\alpha) \sin 2\delta, \end{aligned} \quad (10)$$

$$\begin{aligned} \Delta\delta = & -R_1 \sin \alpha + R_2 \cos \alpha \\ & - D_1 \cos \alpha \sin \delta - D_2 \sin \alpha \sin \delta + D_3 \cos \delta \\ & + E_{20} \sin 2\delta \\ & - (E_{21}^{\text{Re}} \cos \alpha - E_{21}^{\text{Im}} \sin \alpha) \cos 2\delta \\ & - (M_{21}^{\text{Re}} \sin \alpha + M_{21}^{\text{Im}} \cos \alpha) \sin \delta \\ & - (E_{22}^{\text{Re}} \cos 2\alpha - E_{22}^{\text{Im}} \sin 2\alpha) \sin 2\delta \\ & + 2(M_{22}^{\text{Re}} \sin 2\alpha + M_{22}^{\text{Im}} \cos 2\alpha) \cos \delta, \end{aligned} \quad (11)$$

where the  $R_i$  parameters denote the rotations, the  $D_i$  parameters denote the glide terms, and the  $E_{2i}$  and  $M_{2i}$  parameters denote the poloidal (or electric-type) and toroidal (or magnetic-type) quadrupole terms (each of which with a real and an imaginary part). For every pair of catalogs that we have compared (see the following subsections), the 16 parameters of the transformation were obtained by a least-squares fit to the coordinate differences for the common sources. In this fit, the coordinate differences were weighted by the inverse of the sum of the squared coordinate uncertainties in the two catalogs. In order not to bias the determination, the fit was performed after excluding the outliers. As a criterion, a source was considered as an outlier if the angular separation between the measured positions in the two catalogs normalized by its formal uncertainty (a quantity hereafter referred to as normalized separation) is larger than 5. Additionally, all sources with an angular separation larger than 5 mas or with an error ellipse in position with a semi-major axis larger than 5 mas in either catalog were also discarded.

### 6.2. Comparison with ICRF2

The coordinate transformation described in the previous section was first applied to comparing the ICRF3  $S/X$  band frame with its predecessor, ICRF2, also constructed at  $S/X$  band. Since the purpose of the analysis is to assess the deformations between the two frames and not only to check their alignment, the comparison was carried out based on all sources common to the two catalogs and not just the defining sources. After elimination of the outliers, there were a total of 2918 such sources. To facilitate the evaluation of the results, the 16 parameters derived from the fit (i.e., three rotations, three glide parameters, and ten quadrupole terms) are plotted in the form of a bar chart in Fig. 15. As expected, the rotations, which reach  $15 \mu\text{as}$  at most, are small. These cannot be exactly zero because the no-net-rotation constraints imposed to align ICRF3 onto ICRF2 were applied only to the ICRF2 defining sources (see Sect. 3.3). Additionally, the meaning of that alignment in the (new) context of inclusion of Galactic acceleration in the modeling remains somewhat uncertain since ICRF2, unlike ICRF3, did not have a reference epoch. Looking at the other parameters, it is striking that the glide terms stand out, whereas all quadrupole terms but one are not significant. The  $D_2$  and  $D_3$  glide terms show values of  $-63 \pm 4 \mu\text{as}$  and

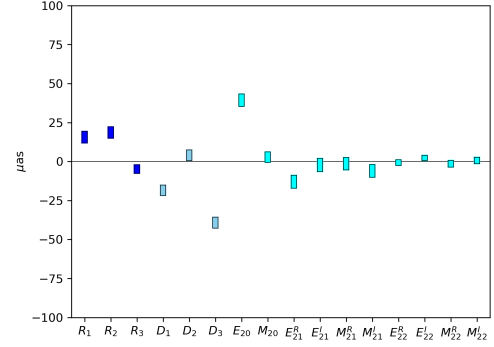


**Fig. 15.** Bar chart showing the values of the 16 parameters of the transformation between the ICRF3 *S/X* band frame and ICRF2. The  $R_i$  parameters are for the rotations, the  $D_i$  parameters are for the glide terms, and the  $E_{2i}$  and  $M_{2i}$  parameters are for the quadrupole terms. See Eqs. (10) and (11) for further details on the transformation.

$-90 \pm 4 \mu\text{as}$ , respectively, well above the expected deformations between the two frames. The only significant quadrupole parameter is the  $E_{20}$  term which shows a value of  $43 \pm 4 \mu\text{as}$ . All the other quadrupole terms are found to be no larger than  $10 \mu\text{as}$ .

In order to try to get insights into such systematics, we produced several variants of the ICRF3 *S/X* band frame by changing the reference epoch of the catalog or alternately by not considering Galactic acceleration in the modeling. Interestingly, the  $D_2$  and  $D_3$  glide terms for these variants were found to vary by several tens of microarcseconds in the comparison to ICRF2, in line with the level of the systematics observed for those terms. Such findings are not unexpected since Galactic acceleration manifests itself as a dipolar deformation in the source coordinates (e.g., Titov & Lambert 2013). Moving further, and noting that Mignard et al. (2016) mentioned this phenomenon as a possibility for explaining the observed glide between ICRF2 and the *Gaia* Data Release 1 (*Gaia* DR1) auxiliary quasar solution, we decided to reproduce an equivalent of ICRF2 by considering only the stretch of data used for ICRF2 (i.e., including only the VLBI sessions up to March 2009 in the solution) and to make a variant that adds Galactic acceleration in the modeling, as implemented for ICRF3. To guarantee the maximum consistency, those two analyses were conducted by employing the same software package as that used for ICRF2, namely CALC-SOLVE (see Fey et al. 2015). Looking at the results, we first observed that our “reproduced” ICRF2 shows similar deformations as the original ICRF2 when compared to ICRF3, hence ruling out the possibility that ICRF2 was in error. Most importantly, the ICRF2 variant that incorporates Galactic acceleration modeling was found to have much reduced glide terms compared to the original or reproduced ICRF2. This is illustrated by the bar chart in Fig. 16 which shows that the  $D_2$  term has now vanished while the  $D_3$  term has been cut by more than half (down to a value of  $-39 \pm 4 \mu\text{as}$ ) in this variant, hence indicating that the deformations between the two frames, in large part, stem from Galactic acceleration not accounted for in ICRF2. This is somehow not surprising since the data set for ICRF2 already covered 30 years, enough for Galactic acceleration effects to emerge, even though the accuracy of the frame was lower than that of ICRF3.

The transformation parameters for the above comparisons are reported in full in Table 13 below. Also included in that table are the values of the parameters derived for the additional comparisons that we have accomplished (i.e., relative to the *Gaia*-CRF2 frame and for the *K* band and *X/Ka* band frames), the details of which are presented in the following subsections.



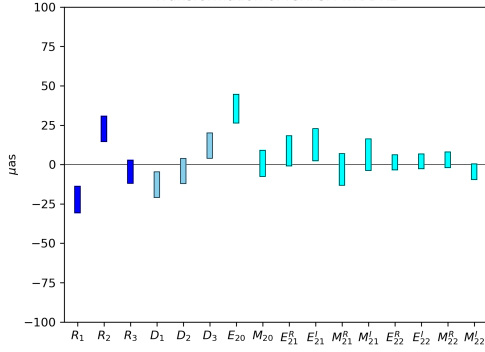
**Fig. 16.** Bar chart showing the values of the 16 parameters of the transformation between the ICRF3 *S/X* band frame and the reproduced ICRF2, the latter incorporating Galactic acceleration in the modeling. The  $R_i$  parameters are for the rotations, the  $D_i$  parameters are for the glide terms, and the  $E_{2i}$  and  $M_{2i}$  parameters are for the quadrupole terms. See Eqs. (10) and (11) for further details on the transformation.

### 6.3. Comparison with *Gaia*-CRF2

The ICRF3 *S/X* band frame was compared in a second stage with the *Gaia*-CRF2 frame, a fully independent frame constructed in the optical domain. For this comparison, a pre-requisite was to identify the sample of sources common to the two frames. While Gaia Collaboration (2018) made available such a sample, we found it necessary to make an update since the original cross-identification of the sources was based on a prototype version of ICRF3 which differs from ICRF3, the final frame not being available by the time of publication. For our determination, we used the same criteria as those originally employed to identify the sources in common with the ICRF3 prototype (Lindgren et al. 2018). Accordingly, a radius of  $0.1''$  was used for the positional matching, supplemented by *Gaia*-specific conditions whose purpose was to reduce the risk of contamination of the quasar sample by Galactic stars. These conditions ensured (i) that a minimum number of eight field-of-view transits was used for each source, (ii) that the source astrometric parameters were obtained exclusively from five-parameter solutions (i.e., including position, parallax, and proper motion), and (iii) that the normalized parallaxes and proper motions are less than 5 (see Lindgren et al. 2018, Eq. (14)). Applying this scheme, a total of 3373 common sources were identified from the positional matching. The additional *Gaia*-specific conditions imposed led to discarding 390 of these (90 from condition (i), 275 from condition (ii), and 25 from condition (iii)), hence leaving a total of 2983 common sources. This is about 6% larger than the number of such sources found by Gaia Collaboration (2018) when comparing *Gaia*-CRF2 to the ICRF3 prototype (2983 vs 2820 common sources).

The transformation parameters between the two frames were determined by assuming that the epochs of the two frames are the same. Though not identical, the ICRF3 and *Gaia*-CRF2 epochs are indeed very close, 2015.0 for ICRF3 and 2015.5 for *Gaia*-CRF2 (mean epoch of the observations). As a result, the  $5.8 \mu\text{as yr}^{-1}$  correction for Galactic acceleration would induce changes in the ICRF3 positions by  $3 \mu\text{as}$  at most, equivalent to one tenth of the ICRF3 noise floor, a value not regarded as significant. After elimination of the outliers, a total of 2612 sources (out of the initial 2983 common sources) was left for the comparison, corresponding to 87.5% of the original sample. This means that 12.5% of the common sources show significantly discrepant radio and optical positions according to the criteria set above, a percentage comparable to that reported by Gaia Collaboration (2018). The results of the fit for the 16 parameters of the





**Fig. 17.** Bar chart showing the values of the 16 parameters of the transformation between the ICRF3 *S/X* band frame and the *Gaia*-CRF2 frame. The  $R_i$  parameters are for the rotations, the  $D_i$  parameters are for the glide terms, and the  $E_{2i}$  and  $M_{2i}$  parameters are for the quadrupole terms. See Eqs. (10) and (11) for further details on the transformation.

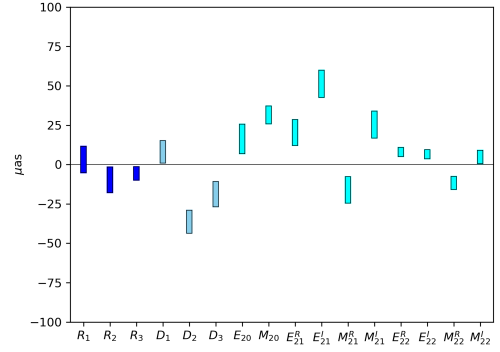
transformation are plotted as a bar chart in Fig. 17, in the same way as previously. As further information, the corresponding parameter values are reported in Table 13. A visual inspection of the bar chart in Fig. 17 reveals in the first place that the two frames are very close. In particular, there is no such a dipolar deformation as that seen for the comparison to ICRF2. The rotation parameters are within  $25 \mu\text{as}$  and have only marginal significance (at most  $2.8\sigma$ , see the error values in Table 13), which means that the two frames are reasonably well aligned. The dipolar terms are less than  $15 \mu\text{as}$  and are non-significant, as are also all of the quadrupole terms but one. The only possibly significant such term is  $E_{20}$  (as for the comparison with ICRF2), which has a value of  $35 \pm 9 \mu\text{as}$  (corresponding to  $3.9\sigma$ ). This term translates into a zonal deformation of the said value, peaking at  $\pm 45^\circ$  declination (with opposite values in the north and in the south) and vanishing at  $0$  and  $\pm 90^\circ$  declination. Such a deformation, if real and inherent to ICRF3, might come from the asymmetry of the VLBI networks which include many more east-west baselines than north-south baselines, as reflected by Fig. 1.

Based on the above comparison, the ICRF3 *S/X* band frame and *Gaia*-CRF2 frame were found to be consistent at the  $30 \mu\text{as}$  level, in agreement with the ICRF3 noise floor, which is also  $30 \mu\text{as}$ . This contrasts with the comparison to ICRF2 discussed in the previous subsection which revealed relative deformations up to about  $100 \mu\text{as}$  between the two frames. Considering that ICRF3 and *Gaia*-CRF2 are fully independent frames, it is likely then that the observed deformations are inherent to ICRF2.

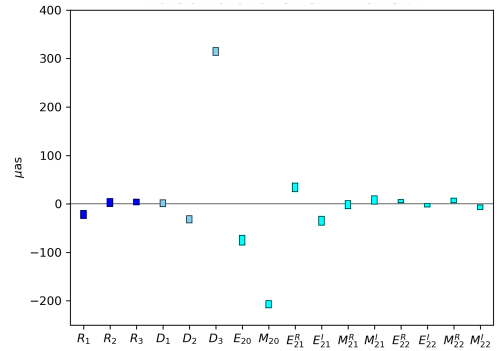
#### 6.4. Intercomparison of the three individual catalogs

In addition to the comparisons relative to the ICRF2 and *Gaia*-CRF2 frames, we also performed internal comparisons between the three independent catalogs that form ICRF3. For this purpose, the *S/X* band catalog was used as a reference against which the *K* band and *X/Ka* band catalogs were compared. Employing the same scheme as before for the outlier elimination, a total of 768 sources was left for the comparison with the *S/X* band catalog at *K* band, corresponding to 97% of the sources common to the two catalogs. At *X/Ka* band, the percentage of sources remaining after such filtering was lower, with only 87% of the common sources (i.e., 556 sources) left for the comparison.

The bar chart in Fig. 18 shows the level of the 16 transformation parameters derived from fitting coordinate differences



**Fig. 18.** Bar chart showing the values of the 16 parameters of the transformation between the ICRF3 *S/X* band and *K* band frames. The  $R_i$  parameters are for the rotations, the  $D_i$  parameters are for the glide terms, and the  $E_{2i}$  and  $M_{2i}$  parameters are for the quadrupole terms. See Eqs. (10) and (11) for further details on the transformation.



**Fig. 19.** Bar chart showing the values of the 16 parameters of the transformation between the ICRF3 *S/X* band and *X/Ka* band frames. The  $R_i$  parameters are for the rotations, the  $D_i$  parameters are for the glide terms, and the  $E_{2i}$  and  $M_{2i}$  parameters are for the quadrupole terms. It should be emphasized that the scale is different from that in Figs. 15–18. See Eqs. (10) and (11) for further details on the transformation.

between the *K* band and *S/X* band catalogs, while that in Fig. 19 shows the equivalent for the *X/Ka* band catalog. Looking at the rotations, these charts confirm that the *K* band and *X/Ka* band catalogs are properly aligned onto the *S/X* band frame, a property to be expected as a result of the analysis configuration used to build those catalogs (see Sect. 3.3). With a single exception, all rotation values are lower than  $10 \mu\text{as}$  and have a significance at the level of one sigma or less (see transformation parameters in Table 13). As already remarked when addressing the alignment to ICRF2, such rotations cannot be exactly zero because the no-net-rotation constraints imposed to align the *K* band and *X/Ka* band frames onto the *S/X* band frame were applied solely to the ICRF3 defining sources. The only rotation parameter departing from zero in these comparisons is the rotation  $R_1$  for the *X/Ka* band catalog comparison, which has a value of  $-22 \pm 8 \mu\text{as}$ , hence with a marginal significance of  $2.7\sigma$ . Unlike rotation parameters, the deformation parameters reveal a different picture at *K* band and *X/Ka* band. At *K* band, all dipole and quadrupole terms are within  $50 \mu\text{as}$ . The largest such terms are the dipole term  $D_2$  ( $-36 \pm 7 \mu\text{as}$ ) and the quadrupole terms  $M_{20}$  ( $31 \pm 6 \mu\text{as}$ ) and  $E_{21}^{\text{lm}}$  ( $51 \pm 9 \mu\text{as}$ ), each of which showing roughly a  $5\sigma$  significance (see Table 13). Considering the weakness of the observing in the far south, such moderate deformations of the *K* band frame are not unexpected. It is also worth noting that this level of deformation is consistent with the noise

**Table 13.** Parameters of the transformations between ICRF3 (*S/X* band catalog) and other frames, including rotation, glide (dipole terms), and quadrupole terms (all in microarcseconds), plus the number of sources involved in the calculations.

Parameters	Reference frame			ICRF3 catalog	
	ICRF2	ICRF2 <sup>GA</sup>	<i>Gaia</i> -CRF2	<i>K</i> band	<i>X/Ka</i> band
Nb sources					
common	3414	3414	2983	793	638
outliers	496	382	371	25	82
used	2918	3032	2612	768	556
Rotation					
$R_1$	$8 \pm 4$	$16 \pm 4$	$-22 \pm 8$	$3 \pm 8$	$-22 \pm 8$
$R_2$	$15 \pm 4$	$19 \pm 4$	$23 \pm 8$	$-10 \pm 8$	$3 \pm 8$
$R_3$	$0 \pm 3$	$-5 \pm 3$	$-5 \pm 7$	$-6 \pm 4$	$4 \pm 5$
Glide (dipole)					
$D_1$	$-22 \pm 4$	$-18 \pm 3$	$-13 \pm 8$	$8 \pm 7$	$1 \pm 7$
$D_2$	$-63 \pm 4$	$4 \pm 3$	$-4 \pm 8$	$-36 \pm 7$	$-32 \pm 7$
$D_3$	$-90 \pm 4$	$-39 \pm 4$	$12 \pm 8$	$-19 \pm 8$	$314 \pm 8$
Quadrupole					
$E_{20}$	$43 \pm 4$	$39 \pm 4$	$35 \pm 9$	$16 \pm 9$	$-75 \pm 10$
$M_{20}$	$5 \pm 4$	$3 \pm 3$	$1 \pm 8$	$31 \pm 6$	$-207 \pm 7$
$E_{21}^{\text{Re}}$	$-11 \pm 5$	$-13 \pm 4$	$9 \pm 10$	$20 \pm 8$	$34 \pm 9$
$E_{21}^{\text{Im}}$	$3 \pm 5$	$-2 \pm 4$	$13 \pm 10$	$51 \pm 9$	$-35 \pm 9$
$M_{21}^{\text{Re}}$	$-1 \pm 4$	$-1 \pm 4$	$-3 \pm 10$	$-16 \pm 8$	$-2 \pm 8$
$M_{21}^{\text{Im}}$	$-5 \pm 4$	$-6 \pm 4$	$6 \pm 10$	$25 \pm 9$	$8 \pm 9$
$E_{22}^{\text{Re}}$	$1 \pm 2$	$-1 \pm 2$	$1 \pm 5$	$8 \pm 3$	$6 \pm 3$
$E_{22}^{\text{Im}}$	$3 \pm 2$	$2 \pm 2$	$2 \pm 5$	$7 \pm 3$	$-3 \pm 4$
$M_{22}^{\text{Re}}$	$-3 \pm 2$	$-1 \pm 2$	$3 \pm 5$	$-12 \pm 4$	$7 \pm 5$
$M_{22}^{\text{Im}}$	$2 \pm 2$	$1 \pm 2$	$-5 \pm 5$	$5 \pm 4$	$-7 \pm 5$

**Notes.** ICRF2<sup>GA</sup> denotes the equivalent of ICRF2 that was reproduced after incorporating Galactic acceleration in the modeling.

floor of the frame ( $30 \mu\text{as}$  in right ascension and  $50 \mu\text{as}$  in declination), as reported in Table 6. In contrast, the bar chart for the *X/Ka* band catalog (Fig. 19) reveals much stronger deformations. In particular, there are three terms that stick out: the glide term  $D_3$  ( $314 \pm 8 \mu\text{as}$ ) and the quadrupole terms  $E_{20}$  ( $-75 \pm 10 \mu\text{as}$ ) and  $M_{20}$  ( $-207 \pm 7 \mu\text{as}$ ). As noted previously, the *X/Ka* band network is comprised of only four sites and thus has an inherent limited observing geometry, which is a possible reason for the deformations seen. Such deformations may also explain the larger number of outliers and the slight misalignment of the frame (reflected by the non-zero value of the  $R_1$  rotation parameter) noted above. In the future, a specific effort should be placed on reducing these systematics by strengthening the observing geometry at this frequency band.

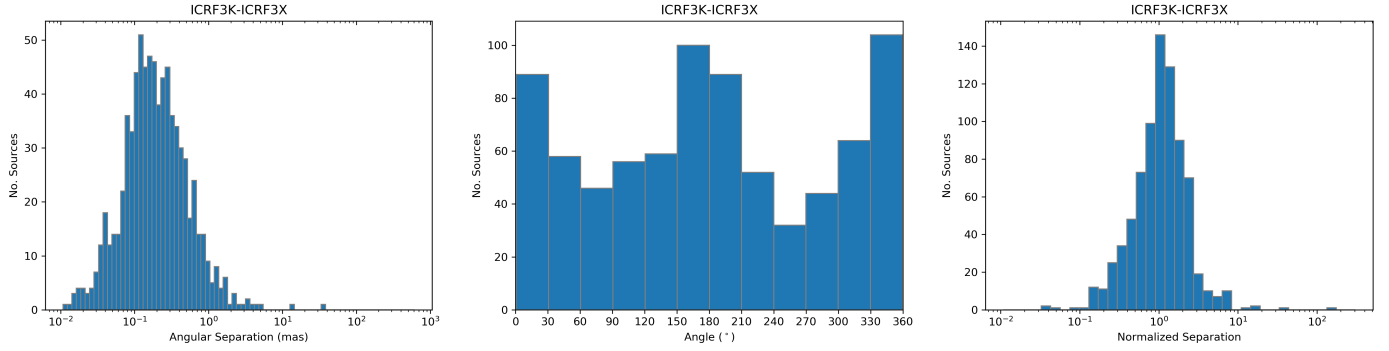
### 6.5. Consistency of source positions at the three frequencies

While the previous sections deal with rotations and deformations between catalogs, it is also of interest to compare the individual source positions measured at the three ICRF3 frequencies. As indicated by the source breakdown in Fig. 5, the *S/X* band catalog has 793 sources in common with the *K* band catalog and 638 sources in common with the *X/Ka* band catalog, which provides adequate material for the desired comparisons. For this purpose, the *S/X* band positions were adopted as a reference against which the *K* band and *X/Ka* band positions were compared. Prior to the comparison of individual source coordinates, the 16-parameter transformations in Table 13 were applied to the

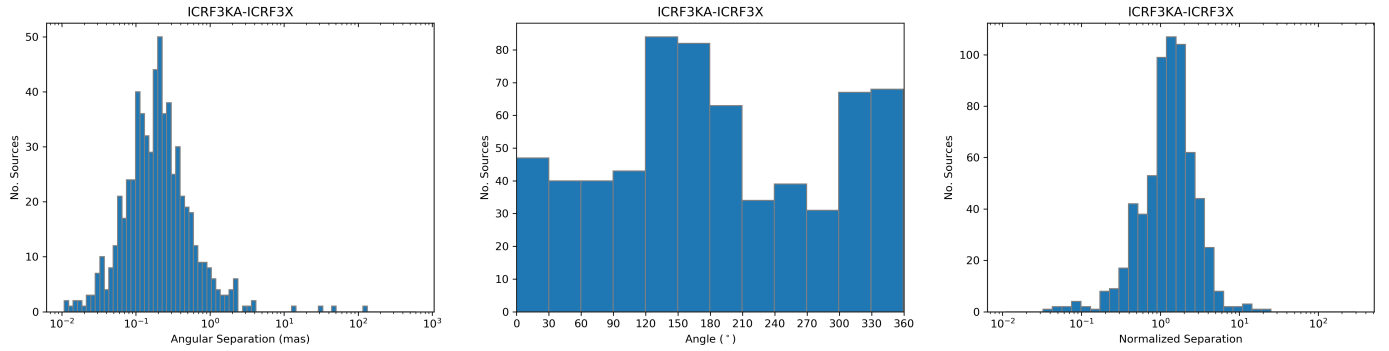
*K* band and *X/Ka* band catalogs in order to free the comparisons from catalog deformations. Aside from source coordinate differences, we examined also the angular separation and direction of the offset vector joining the measured positions at *K* band (or *X/Ka* band) and *S/X* band, along with the corresponding normalized separation. The results of these comparisons are presented graphically in Figs. 20 and 21 for those three quantities.

Looking at the distribution of angular separations (left-hand panels in Figs. 20 and 21), both histograms are found to peak at about 0.15–0.20 mas, which is consistent with the median ellipse error of the catalogs (see common sources in Table 8), though slightly larger. Examining further the *K* band distribution, the first quartile (25% of data) is at 0.10 mas, the second quartile (50% of data) is at 0.18 mas, while the third quartile (75% of data) is at 0.36 mas. Values for the *X/Ka* band distribution are very similar, 0.10 mas for the first quartile, 0.19 mas for the second quartile, and 0.34 mas for the third quartile. As regards offset vector directions, the histograms in Figs. 20 and 21 (middle panels) are indicative of non-uniform distributions. In both cases, the observed directions show an excess at  $0^\circ$  and  $180^\circ$ , that is, along the declination axis. In a separate check, we also compared directly the *K* band and *X/Ka* band positions (plots not shown) and found a similar excess, hence ruling out that it comes purely from the *S/X* band data. Such an excess is not expected from the source physics since any jet-like features that could induce offsets between positions measured at different frequencies should have a random orientation due to the VLBI jets having no preferred direction in the sky. It is thus most likely that the observed excess arises from remaining catalog systematics, especially in declination, not fully removed by the applied transformations.

Looking at the right-hand panels in Figs. 20 and 21, the distribution of normalized separations is found to peak near 1 in the case of the *K* band comparison and around 1.5 in the case of the *X/Ka* band comparison. The median value of the distribution is 1.09 for the former and is 1.32 for the latter. This is an indirect indication that source position uncertainties in all are reasonably well estimated, though possibly somewhat underestimated at *X/Ka* band. A further noticeable feature in these histograms is that a visible portion of the sources show normalized separations above 3. There are 46 sources in that condition (i.e., 6% of the set of common sources), including five ICRF3 defining sources (0403–132, 1245–457, 1448+762, 1745+624, and 2229+695), in the case of the *K* band comparison (see Table 14) and 70 such sources (i.e., 11% of the set of common sources), including 18 ICRF3 defining sources, in the case of the *X/Ka* band comparison (see Table 15). We suspect that the higher percentage found for the latter is due to the uncertainty in the angular separations being underestimated, as reflected by the median normalized separation of 1.32 noted above. Indeed, if assuming a 30% underestimation, hence considering only sources with normalized separations above 3.9, the number of sources that stand out falls down to 32 (i.e., 5% of the set of common sources), including three ICRF3 defining sources (0607–157, 0648–165, and 0743–006). In any case, that percentage of outliers (i.e., about 5% for both the *K* band and *X/Ka* band comparisons) is much higher than anticipated for a Rayleigh distribution (where only 1.1%, i.e., 7–8 sources, would be expected) and it only reflects the existence of systematics in the positions of these sources. It is of interest that a number of sources in those two tables are known to have extended VLBI structures, for example 0430+052 (3C120), 0923+392 (4C39.25), or 2251+158



**Fig. 20.** Comparison of the ICRF3 positions at *S/X* band and *K* band for the 793 sources common to the two catalogs. Position differences were derived after applying the transformation in Table 13 to the *K* band source coordinates. The histogram in the *left-hand panel* shows the distribution of the angular separation between the two sets of positions, that in the *middle panel* shows the distribution of the direction of the offset vector joining those positions (counted counter clockwise), and that in the *right-hand panel* the distribution of the normalized separation.



**Fig. 21.** Comparison of the ICRF3 positions at *S/X* band and *X/Ka* band for the 638 sources common to the two catalogs. Position differences were derived after applying the transformation in Table 13 to the *X/Ka* band source coordinates. The histogram in the *left-hand panel* shows the distribution of the angular separation between the two sets of positions, that in the *middle panel* shows the distribution of the direction of the offset vector joining those positions (counted counter clockwise), and that in the *right-hand panel* the distribution of the normalized separation.

**Table 14.** Positional offsets for the 46 sources for which the normalized separation  $N_L$  between the *S/X* and *K* band positions is above 3.

Source name	Coordinate offsets		Offset vector			$N_L$
	$\Delta \cos \delta$ ( $\mu$ as)	$\Delta \delta$ ( $\mu$ as)	Length <sup>(†)</sup> ( $\mu$ as)	Direction <sup>(‡)</sup> (°)		
0003–066	$-86 \pm 57$	$-423 \pm 111$	$431 \pm 109$	$191 \pm 8$		4.0
0014+813	$-140 \pm 76$	$-438 \pm 95$	$460 \pm 94$	$198 \pm 10$		4.9
0112–017	$-601 \pm 86$	$-62 \pm 149$	$604 \pm 88$	$264 \pm 14$		6.9
0146+056	$589 \pm 77$	$-338 \pm 135$	$679 \pm 95$	$120 \pm 10$		7.2
0212+735	$330 \pm 62$	$-320 \pm 73$	$460 \pm 68$	$134 \pm 8$		6.8
0229+131	$-308 \pm 64$	$57 \pm 107$	$313 \pm 66$	$281 \pm 19$		4.7

**Notes.** The content here is printed only for the first six sources. The table in its entirety is available in electronic form from the CDS. <sup>(†)</sup>The vector length is also denoted as angular separation in the text. <sup>(‡)</sup>The vector direction is counted counter clockwise from north to east.

(3C454.3). Additionally, about half of the sources that are found to deviate in the *K* band comparison (24 sources) happen to deviate also in the *X/Ka* band comparison. Taking a step further, we examined the X band structure indices (from BVID) for these sources. For the 39 deviating sources in Table 14 (*K* band comparison) for which structure was assessed, the structure indices range from 1.6 to 4.9, with a mean value of 3.3, hence reflecting a predominance of significantly-structured sources. The structure indices for the 61 deviating sources in Table 15 (*X/Ka* band comparison) for which structure was assessed cover the same range,

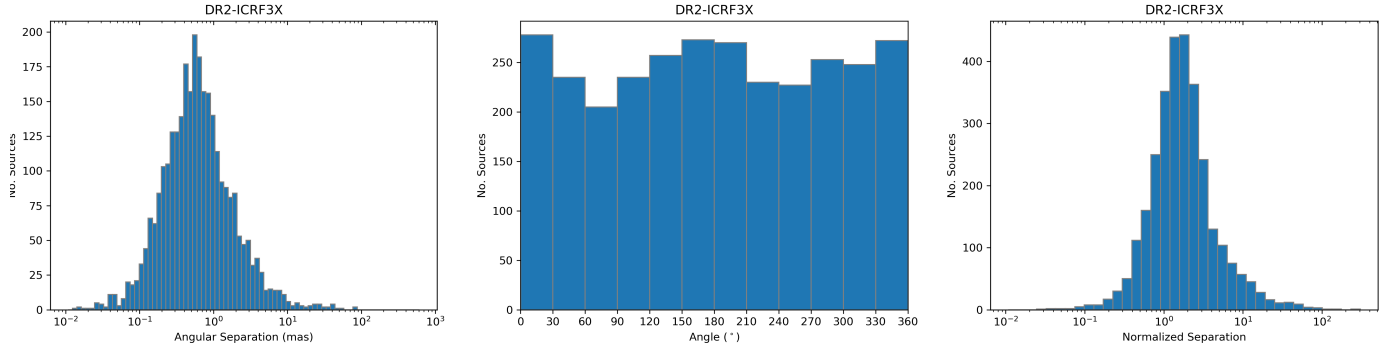
**Table 15.** Positional offsets for the 70 sources for which the normalized separation  $N_L$  between the *S/X* and *X/Ka* band positions is above 3.

Source name	Coordinate offsets		Offset vector			$N_L$
	$\Delta \cos \delta$ ( $\mu$ as)	$\Delta \delta$ ( $\mu$ as)	Length <sup>(†)</sup> ( $\mu$ as)	Direction <sup>(‡)</sup> (°)		
0003–066	$-109 \pm 98$	$544 \pm 133$	$555 \pm 132$	$349 \pm 10$		4.2
0038–020	$-147 \pm 114$	$674 \pm 191$	$690 \pm 188$	$348 \pm 10$		3.7
0059+581	$182 \pm 58$	$-163 \pm 69$	$244 \pm 63$	$132 \pm 15$		3.9
0112–017	$-593 \pm 96$	$116 \pm 131$	$604 \pm 98$	$281 \pm 12$		6.2
0119+115	$-23 \pm 62$	$-324 \pm 93$	$325 \pm 93$	$184 \pm 11$		3.5
0122–003	$-101 \pm 126$	$823 \pm 205$	$829 \pm 204$	$353 \pm 9$		4.1

**Notes.** The content here is printed only for the first six sources. The table in its entirety is available in electronic form from the CDS. <sup>(†)</sup>The vector length is also denoted as angular separation in the text. <sup>(‡)</sup>The vector direction is counted counter clockwise from north to east.

from 1.6 to 4.9, but have a mean value that is somewhat lower (3.0). On the other hand, if reducing the sample to the 32 sources with normalized separations above 3.9 (see the discussion about underestimation of the uncertainties above), the mean value of the structure index goes up to 3.3. The *X/Ka* band comparison therefore also provides indication that the deviating sources have significant structures. In all, those extended morphologies may well explain the observed position inconsistencies between *S/X* and *K* or *X/Ka* band. The further investigation and interpretation of these offsets, however, is beyond the scope of this paper.





**Fig. 22.** Comparison of the ICRF3 positions at *S/X* band with the *Gaia*-CRF2 positions for the 2983 sources common to the two frames. Position differences were derived after applying the transformation in Table 13 to the *Gaia*-CRF2 source coordinates. The histogram in the *left-hand panel* shows the distribution of the angular separation between the two sets of positions, that in the *middle panel* shows the distribution of the direction of the offset vector joining those positions (counted counter clockwise), and that in the *right-hand panel* the distribution of the normalized separation.

### 6.6. Consistency of radio and optical source positions

The individual ICRF3 source positions may also be compared to the *Gaia*-CRF2 optical positions. For this characterization, we used the *S/X* band catalog since it contains the largest number of sources in common (2983 sources, among which 250 defining sources) and adopted the same scheme as above for the comparison, meaning that the transformation in Table 13 was applied to the *Gaia*-CRF2 frame prior to comparing the radio and optical positions. As before, the separations, offset vector directions, and normalized separations between the two sets of positions were computed, and the distribution of these is shown in Fig. 22.

Looking at the distribution of angular separations (left-hand panel in Fig. 22), the histogram is found to peak at about 0.6 mas, with the first quartile at 0.30 mas, the second quartile at 0.58 mas, and the third quartile at 1.13 mas. These values are roughly a factor of three larger than those obtained when comparing the *S/X* band catalog to the *K* band or *X/Ka* band catalog. Such a difference might be explained, at least in part, by the common sources in this comparison having generally larger position uncertainties than those involved in the previous comparisons. This is true both for the *Gaia*-CRF2 positions, where the median value of the ellipse error semi-major axis for the 2983 common sources is 0.26 mas, and the *S/X* band positions, where this quantity is 0.19 mas. In contrast, the corresponding quantities are 0.14 mas and 0.09 mas in the *K* band to *S/X* band comparison (with 793 common sources) and 0.11 mas and 0.07 mas in the *X/Ka* band to *S/X* band comparison (with 638 common sources). The reason why the *S/X* band median position uncertainty is larger when comparing to the *Gaia*-CRF2 catalog than when comparing to the *K* band or *X/Ka* band catalog is that the common sources fall predominantly into the VLBA-only category in that comparison, unlike those involved in the *K* band and *X/Ka* band comparisons where other sources (i.e., observed also by the IVS) predominate. Looking now at the middle panel in Fig. 22, which plots the direction of the offset vectors, the distribution is found to be much more uniform than when comparing to the *K* band and *X/Ka* band catalogs (see Figs. 20 and 21), though there is still a small excess along the declination axis. The presence of this excess indicates that the *S/X* band catalog is not entirely free from declination systematics since there is no reason why such systematics should come from the *Gaia*-CRF2 catalog where declination plays no particular role due to the frame being built from space. The effect,

however, is less pronounced than that observed in the *K* band and *X/Ka* band catalogs.

Of particular interest is the examination of the distribution of normalized separations, which is shown in the right-hand panel of Fig. 22. The histogram peaks between 1 and 2, with the median value of the distribution at 1.64, hence indicating possibly an overall underestimation of the position uncertainties (which for the VLBI part would be in contrast with the findings in the previous subsection) or the presence of real physical offsets between the measured optical and radio positions. Another noticeable feature in that distribution is that a significant portion of the sources show large normalized separations (see the higher end of the histogram). Looking at this feature in detail, a total of 653 sources are found to have a normalized separation larger than 3, including 59 defining sources. Interestingly, the percentage of sources that stand out is similar when considering the entire set of common sources or only the defining sources (about 22% for the former and 24% for the latter). Furthermore, there are 144 sources among these that show normalized separations larger than 10 (corresponding to 4.8% of the common sources), including four defining sources (1.6% of the defining sources in common). All 653 sources in this condition (i.e., with a normalized separation larger than 3) are listed in Table 16, together with the measured coordinate offsets, offset vectors, and normalized separations. The existence of such notable VLBI-*Gaia* position offsets for a significant portion of the sources was first revealed by Mignard et al. (2016), Kovalev et al. (2017), and Petrov & Kovalev (2017b) based on the *Gaia* DR1 catalog (Lindgren et al. 2016) and interpreted as the manifestation of optical jets on scales 1–100 mas (Petrov & Kovalev 2017a). The release of the *Gaia* DR2 catalog (Lindgren et al. 2018) confirmed those findings with a higher level of significance (Gaia Collaboration 2018; Petrov et al. 2019; Plavin et al. 2019). Following the lines of the previous section, we examined the structure indices of the sources with significant positional offsets when available. Out of the 653 sources in Table 16, 267 were found to have an X band structure index in BVID. The mean value of these structure indices is 3.4, with a dichotomy between the 50 defining sources (mean structure index of 2.7) and the 217 other sources (mean structure index of 3.5), hence indicating that sources with significant structure predominate. The ICRF3 data thus also suggest that the large *Gaia*-VLBI offsets observed for a notable fraction of the sources are likely a manifestation of source structure.

**Table 16.** Positional offsets for the 653 sources for which the normalized separation  $N_L$  between the  $S/X$  and *Gaia*-CRF2 positions is above 3.

Source name	Coordinate offsets		Offset vector		$N_L$
	$\Delta \cos \delta$ ( $\mu$ as)	$\Delta \delta$ ( $\mu$ as)	Length <sup>(†)</sup> ( $\mu$ as)	Direction <sup>(‡)</sup> ( $^\circ$ )	
2357–326	$-6265 \pm 177$	$-618 \pm 289$	$6295 \pm 178$	$264 \pm 3$	35.3
2359–221	$-1702 \pm 518$	$1114 \pm 567$	$2034 \pm 533$	$303 \pm 16$	3.8
0000–199	$-1053 \pm 339$	$1584 \pm 359$	$1902 \pm 353$	$326 \pm 10$	5.4
0001–120	$1108 \pm 319$	$50 \pm 328$	$1109 \pm 319$	$87 \pm 17$	3.5
0003+380	$6425 \pm 233$	$-5409 \pm 185$	$8399 \pm 214$	$130 \pm 1$	39.2
0003–066	$-47 \pm 99$	$227 \pm 70$	$232 \pm 71$	$348 \pm 24$	3.3

**Notes.** The content here is printed only for the first six sources. The table in its entirety is available in electronic form from the CDS. <sup>(†)</sup>The vector length is also denoted as angular separation in the text. <sup>(‡)</sup>The vector direction is counted counter clockwise from north to east.

## 7. Adoption of ICRF3 by the IAU

As stated in the mandate of the working group, ICRF3 was presented at the 30th IAU General Assembly held in Vienna (Austria) on August 20–31, 2018. The resolution adopted by the General Assembly, referenced as Resolution B2<sup>11</sup>, resolved that the fundamental ICRS realization shall be ICRF3 from January 1, 2019, that the organizations responsible for astrometric and geodetic VLBI observing programs (e.g., IVS) shall take appropriate measures to both maintain and improve ICRF3, and that the organizations responsible for defining high-accuracy reference frames at other wavelengths, together with the IERS, shall take appropriate measures to align those reference frames onto ICRF3 with the highest possible accuracy. As a result of this adoption, ICRF3 replaced ICRF2 as the fundamental celestial reference frame to use for all applications on January 1, 2019.

## 8. Evolution of the ICRF

The material presented in the previous sections already provides directions at which to target prospective VLBI observing along with modeling improvements and refinements of the analysis configuration in order to enhance the ICRF in the future. Each of these aspects is discussed in turn in the following paragraphs.

On the observing side, a special effort should be made in the first place to strengthen the celestial frame in the far south. As noted above, the  $S/X$  band frame shows a deficit of sources at the most southern declinations (Fig. 6). While the number of sources is similar in the south and in the north for declinations between  $-35^\circ$  and  $+35^\circ$  (1461 vs 1460 sources), the number of sources further south ( $< -35^\circ$  declination) is lower by a factor of 2.5 than that further north ( $> 35^\circ$  declination) (460 vs 1155 sources). The same applies to the  $K$  band frame, although the deficit of sources in the south is only by a factor of 1.5 in this case. Additionally, it is found that the source position accuracy deteriorates with declination (Fig. 9). For such reasons, it is desirable that a proper observing strategy be devised for these two frequency bands to both increase the source density and augment the source position accuracy in the south. Efforts in this direction have already been engaged and IVS sessions in the south are now strengthened (de Witt et al. 2019). At  $X/Ka$  band, the frame is more uni-

**Table 17.** ICRF3 defining sources for which VLBI structure was not assessed prior to their selection as defining sources.

Source names				
0009–148	0044–846	0227+403	0642–349	0742–562
0802–010	0804–267	0841–607	0926–039	0930–080
1016–311	1036–529	1101–536	1245–457	1312–533
1325–558	1412–368	1511–558	1556–245	1606–398
1753+204	2037+216	2111+400	2121+547	2220–351

form but suffers from systematics. The focus of the observing in the future should then be placed on reducing these systematics.

A second direction for prospective VLBI observing lies in the further increase in the source position accuracy for the bulk of the sources in the  $S/X$  band frame. Figure 6 shows that there are two peaks in the distribution of position uncertainties, a primary peak at  $100\text{--}200\mu\text{as}$  capturing the majority of the sources and a secondary peak just above the noise floor ( $30\mu\text{as}$ ) that captures the roughly 500 sources that have the most precise positions. As noted above, the primary peak (already existing in ICRF2) was brought down from  $350\text{--}700\mu\text{as}$  in ICRF2 to the present error level thanks to a number of VLBA campaigns conducted since 2014 (see Sect. 2.1). Acquiring more data on the corresponding such 3000 sources will bring this peak further down and even closer to the secondary peak at the noise floor. A factor of 3–4 (the same as that gained from ICRF2 to ICRF3) is still to be gained to merge the two peaks. In this respect, continuation of the VLBA campaigns will be decisive to achieve this goal. In this process, observation of the optically bright ICRF3 sources (i.e., detected by *Gaia*) will be also of special interest to strengthen the alignment between the radio and optical frames.

Above all, it will be essential to monitor closely the defining sources. For this purpose, it would be desirable that these be incorporated into the regular geodetic VLBI observing programs, such as those carried out by the IVS, in order to control their astrometric stability. Additionally, it will be important that these sources be imaged on a regular basis to track potential source structure changes. This especially applies to the 25 ICRF3 defining sources for which we had no VLBI images in hand – the list of which is given in Table 17 – so that the brightness distribution of those sources, and hence their suitability as defining sources, may be assessed. A specific effort should also be made toward observing the ICRF3 defining sources at  $K$  band and  $X/Ka$  band since not all of them are included in the respective catalogs and the amount of data for those that do remains too limited in some cases. Having such a plan is necessary to strengthen the alignment between the three catalogs in future realizations of the ICRF. Beyond that, prospective VLBI observing should also aim at filling in the 21 empty sectors, which means to search and identify a suitable defining source for every such sector. These sectors are identified in Table 18 based on the corresponding right ascension and declination range. The table also lists the sources originally selected in these sectors but then discarded due to their having too extended (category C) structures (see Sect. 5.2). In the longer term, one should also seek to replace the 62 defining sources that show moderately extended (i.e. category B) structures by new ones which are more point-like. Such sources are listed in Table 19 along with the identification of the corresponding sectors where new sources are to be found. For the goal of finding new sources, either to fill in

<sup>11</sup> See Resolution B2 at [https://iau.org/static/resolutions/IAU2018\\_ResolB2\\_English.pdf](https://iau.org/static/resolutions/IAU2018_ResolB2_English.pdf).

**Table 18.** Identification of the 21 sectors over the celestial sphere where no suitable defining source was found, either because there is no ICRF3 source in that sector or because the source available (indicated in parentheses) has poor (category C) structure.

Sector number	Right Asc.		Declination		Source name
	Min (h m)	Max (h m)	Min (°)	Max (°)	
24	01 20	02 40	-26.39	-19.47	(0135-247)
26	01 20	02 40	-12.84	-6.38	(0138-097)
35	01 20	02 40	51.06	62.73	(0144+584)
44	02 40	04 00	-12.84	-6.38	(0238-084)
50	02 40	04 00	26.39	33.75	(0333+321)
70	04 00	05 20	41.81	51.06	(0420+417)
126	08 00	09 20	62.73	90.00	(0836+710)
143	09 20	10 40	51.06	62.73	(0917+624)
161	10 40	12 00	51.06	62.73	(1038+528)
171	12 00	13 20	-6.38	0.00	(1253-055)
178	12 00	13 20	41.81	51.06	(1216+487)
230	16 00	17 20	26.39	33.75	(1600+335)
239	17 20	18 40	-33.75	-26.39	
263	18 40	20 00	6.38	12.84	(1947+079)
269	18 40	20 00	51.06	62.73	(1954+513)
272	20 00	21 20	-62.73	-51.06	
275	20 00	21 20	-33.75	-26.39	(2000-330)
287	20 00	21 20	51.06	62.73	(2037+511)
304	21 20	22 40	41.81	51.06	(2200+420)
309	22 40	24 00	-51.06	-41.81	(2326-477)
318	22 40	24 00	12.84	19.47	(2251+158)

**Notes.** Each sector is numbered and defined by a range in right ascension and a range in declination.

the empty sectors in Table 18 or to replace the category B sources in Table 19, the VLBA, again, should be important since it has the capability to observe weaker sources than geodetic VLBI networks.

Aside from acquiring further VLBI data along the lines described above, the evolution of the frame will also depend at some level on the refinements of the modeling. One area where improvement is foreseen relates to the determination of the solar system acceleration vector. As the data accumulate, the time span will also extend, leading to an estimate of the vector amplitude with ever increased accuracy. Future data should also reveal whether that vector is directed entirely toward the Galactic center or is subject to some offset. Additionally, the *Gaia* mission is expected to deliver estimates of those parameters in future data releases, which will further contribute to improving such determination. Another question to be tackled relates to the source structure, which manifests itself through systematics in the VLBI delay measurements (Xu et al. 2019) and instabilities in the individual source positions (Gattano et al. 2018). In the future, those effects will become more prominent as the precision of the measurements continues to increase, making it necessary to consider them in the modeling to further increase the quality of the frame. While the corresponding theoretical framework has been put forward long ago (Charlot 1990), the practical implementation has not been straightforward due to the requirement to have multi-epoch images available for all the sources. While producing such series of images is an enormous task, the advent of several image databases (Collioud & Charlot 2019; Hunt et al. 2019) now offers interesting prospects for the coming years in this area.

Finally, another route to investigate that might possibly help to enhance the frame lies in exploring alternate analysis

**Table 19.** Identification of the 62 sectors over the celestial sphere where the related ICRF3 defining sources, also specified in the table, show moderately extended (category B) structures.

Sector number	Right Asc.		Declination		Defining source
	Min (h m)	Max (h m)	Min (°)	Max (°)	
2	00 00	01 20	-62.73	-51.06	0047-579
14	00 00	01 20	26.39	33.75	0046+316
16	00 00	01 20	41.81	51.06	0110+495
19	01 20	02 40	-90.00	-62.73	0230-790
31	01 20	02 40	19.47	26.39	0149+218
36	01 20	02 40	62.73	90.00	0159+723
46	02 40	04 00	0.00	6.38	0305+039
53	02 40	04 00	51.06	62.73	0302+625
54	02 40	04 00	62.73	90.00	0346+800
59	04 00	05 20	-33.75	-26.39	0400-319
66	04 00	05 20	12.84	19.47	0507+179
72	04 00	05 20	62.73	90.00	0454+844
73	05 20	06 40	-90.00	-62.73	0530-727
80	05 20	06 40	-12.84	-6.38	0605-085
81	05 20	06 40	-6.38	0.00	0539-057
90	05 20	06 40	62.73	90.00	0615+820
91	06 40	08 00	-90.00	-62.73	0738-674
99	06 40	08 00	-6.38	0.00	0743-006
100	06 40	08 00	0.00	6.38	0736+017
101	06 40	08 00	6.38	12.84	0748+126
105	06 40	08 00	33.75	41.81	0641+392
107	06 40	08 00	51.06	62.73	0749+540
109	08 00	09 20	-90.00	-62.73	0842-754
111	08 00	09 20	-51.06	-41.81	0809-493
112	08 00	09 20	-41.81	-33.75	0826-373
115	08 00	09 20	-19.47	-12.84	0818-128
121	08 00	09 20	19.47	26.39	0834+250
125	08 00	09 00	51.06	62.73	0800+618
127	09 20	10 40	-90.00	-62.73	1022-665
133	09 00	10 40	-19.47	-12.84	1027-186
139	09 00	10 40	19.47	26.39	1012+232
150	10 40	12 00	-26.39	-19.47	1143-245
160	10 40	12 00	41.81	51.06	1150+497
173	12 00	13 20	6.38	12.84	1236+077
185	13 20	14 40	-33.75	-26.39	1406-267
186	13 20	14 40	-26.39	-19.47	1435-218
197	13 20	14 40	51.06	62.73	1418+546
199	14 40	16 00	-90.00	-62.73	1448-648
202	14 40	16 00	-41.81	-33.75	1451-400
205	14 40	16 00	-19.47	-12.84	1443-162
206	14 40	16 00	-12.84	-6.38	1510-089
209	14 40	16 00	6.38	12.84	1502+106
213	14 40	16 00	33.75	41.81	1504+377
216	14 40	16 00	62.73	90.00	1448+762
228	16 00	17 20	12.84	19.47	1717+178
233	16 00	17 20	51.06	62.73	1623+578
234	16 00	17 20	62.73	90.00	1642+690
241	17 20	18 40	-19.47	-12.84	1730-130
244	17 20	18 40	0.00	6.38	1725+044
253	18 40	20 00	-90.00	-62.73	1935-692
254	18 40	20 00	-62.73	-51.06	1925-610
256	18 40	20 00	-41.81	-33.75	1954-388
257	18 40	20 00	-33.75	-26.39	1921-293
260	18 40	20 00	-12.84	-6.38	1937-101
276	20 00	21 20	-26.39	-19.47	2037-253
284	20 00	21 20	26.39	33.75	2113+293
289	21 20	22 40	-90.00	-62.73	2142-758
294	21 20	22 40	-26.39	-19.47	2210-257
297	21 20	22 40	-6.38	0.00	2216-038
306	21 20	22 40	62.73	90.00	2229+695
312	22 40	24 00	-26.39	-19.47	2331-240
315	22 40	24 00	-6.38	0.00	2335-027

**Notes.** Each sector is numbered and defined by a range in right ascension and a range in declination.



configurations. While the three catalogs that form ICRF3 have been derived independently, the corresponding data sets could be processed together. The observing system is largely similar at the three frequency bands and such a joint analysis should benefit from the strengths of each data set, most likely resulting in improvements of the overall frame. For example, systematics in the  $X/Ka$  band catalog (see Sect. 6.4) might end up reduced after incorporation of the  $S/X$  band and  $K$  band data sets in the analysis. There are also reasons to expect an even closer alignment of the three catalogs with this scheme. It should be emphasized, though, that the source positions should be kept as separate parameters at the three frequencies, as otherwise any actual physical offsets, such as those suspected here for a fraction of the sources (see Sect. 6.5), would affect the resulting source position estimates. It is worth noting that such a multi-frequency analysis has been recently tried, though not through using a unique software package but instead through a combination of normal equations derived at each frequency band, demonstrating that there would be some value in investigating further such a combination for future realizations of the ICRF (Karbon & Nothnagel 2019). Going further, one may even consider incorporating data from other space geodetic techniques in the combination. Those techniques include satellite and lunar laser ranging, global navigation satellite systems, and the “doppler orbitography and radiopositioning integrated by satellite” system, which, together with VLBI, contribute to the realization of the ITRF. While the impact on the celestial frame may be limited, such a multi-technique combination would have the advantage to deliver unified terrestrial and celestial frames (Seitz et al. 2014).

## 9. Conclusion

A new realization of the ICRF, denoted as ICRF3, has been produced based on VLBI data acquired over nearly 40 years in three frequency bands ( $S/X$ ,  $K$ , and  $X/Ka$  band). This new realization is the first multi-frequency celestial reference frame ever generated. Another new feature incorporated in ICRF3 is the modeling of the galactocentric acceleration of the solar system, which considers an acceleration vector pointing toward the Galactic center with an amplitude of  $5.8 \mu\text{as yr}^{-1}$ . The latter was derived by an adjustment directly to the  $S/X$  band data. Source positions are reported for epoch 2015.0 and the above amplitude value should be used to propagate those positions to other epochs if necessary. The new frame includes positions for 4536 sources at  $S/X$  band, 824 sources at  $K$  band, and 678 sources at  $X/Ka$  band, for a total number of 4588 sources, where 600 of these sources have independent positions available at the three frequencies.

The noise floor in the individual source coordinates is at the level of  $30 \mu\text{as}$ . At  $S/X$  band, the median uncertainty is  $127 \mu\text{as}$  in right ascension and  $218 \mu\text{as}$  in declination, more than a factor of three improvement over the previous realization, ICRF2. This improvement reflects the efforts that have been accomplished during the past decade to re-observe all VCS-type sources (as identified in ICRF2) with the VLBA, while taking the opportunity to augment the number of observed sources by one-third at the same time. Most notably, the  $S/X$  band catalog includes also a pool of 500 sources, observed as part of IVS programs, which have highly accurate positions, with uncertainties in the range of  $30\text{--}60 \mu\text{as}$ . The positional accuracy at  $K$  band and  $X/Ka$  band approaches that at  $S/X$  band but remains a factor of  $1.5\text{--}2$  lower. A subset of 303 sources among the most observed ones at  $S/X$  band has been identified for defining the frame based

on their sky distribution, position stability, and the amount of structure.

ICRF3 is aligned onto ICRF2 to the accuracy of the latter. Comparing the  $S/X$  band frame with the recently released *Gaia*-CRF2 frame in the optical domain, the two frames show no relative deformations above  $30 \mu\text{as}$ . On the other hand, ICRF2 is found to have a significant dipolar deformation, approaching  $100 \mu\text{as}$ , with respect to ICRF3, an effect that in large part results from not considering Galactic acceleration in the modeling for ICRF2. The  $K$  band catalog shows no deformations above  $50 \mu\text{as}$  with respect to the  $S/X$  band frame, unlike the  $X/Ka$  band catalog which suffers from dipolar and quadrupolar systematics. The latter is likely due to the limited geometry of the  $X/Ka$  band data set. Comparisons of the individual source positions at the three frequency bands reveal significant offsets between frequencies for about 5% of the sources, a percentage that increases to 22% when comparing the ICRF3 positions to the *Gaia*-CRF2 optical positions. There are indications that those positional offsets may be the manifestation of extended source structures.

ICRF3 was adopted by the IAU during its 30th General Assembly held in Vienna in August 2018 and is now the fundamental celestial reference frame to be used for all applications. Looking into the future, this first multi-frequency ICRF realization may be regarded as a first stage toward a fully integrated multi-band frame incorporating also the optical *Gaia* data.

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